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G.A. 754496

Stefano Gariazzo

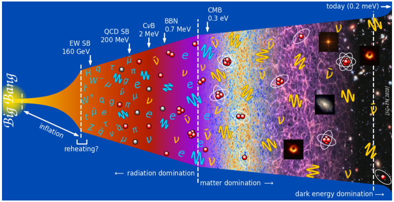
INFN, Turin section
Turin (IT)

gariazzo@to.infn.it
http://personalpages.to.infn.it/~gariazzo/

Active and sterile neutrinos in the early universe: precision calculations

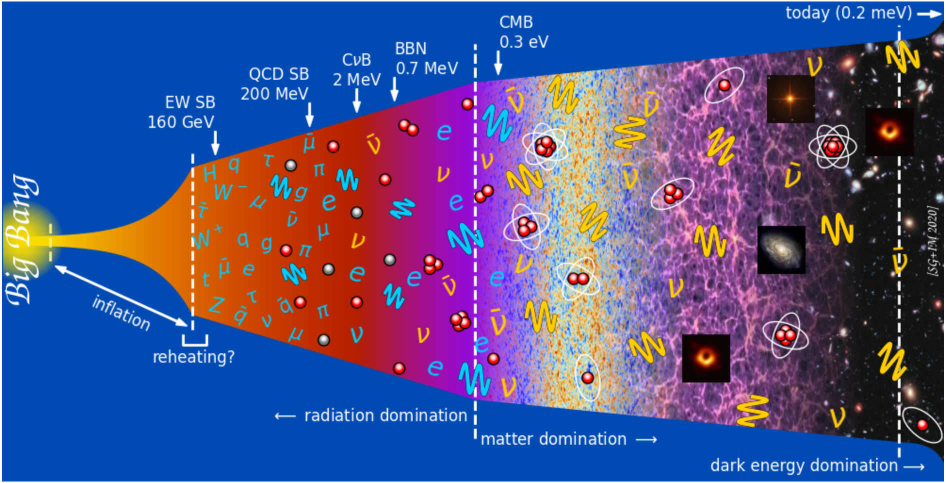
Mostly based on arxiv:2012.02726,
JCAP 07 (2019) 014, arxiv:2003.02289

General annual meeting of the Fellini programme, online, 04/03/2021



today (0.2 meV)
CMB
0.3 eV
BBN
0.7 MeV
CνB
2 MeV
QCD SB
200 MeV
EW SB
160 GeV
Big Bang
inflation
reheating?
⟵ radiation domination
matter domination ⟶
dark energy domination ⟶

History of the universe



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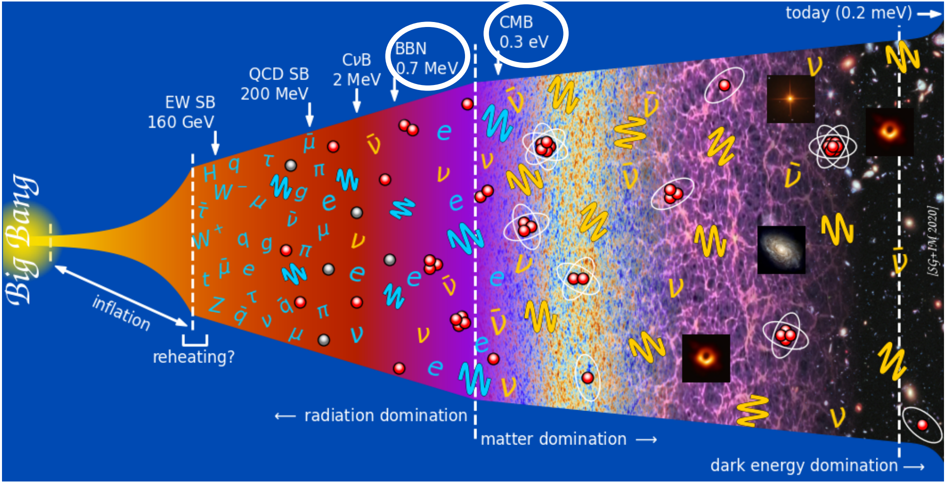

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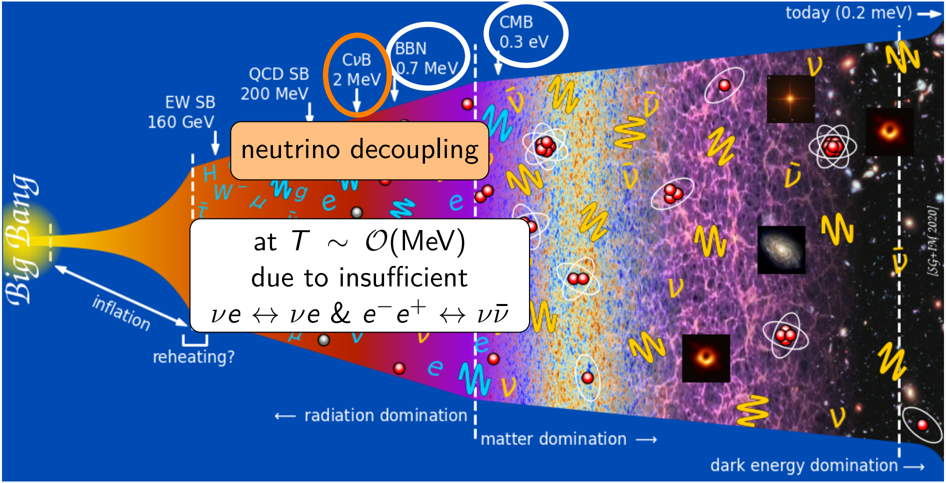


neutrino decoupling

at $T \sim \mathcal{O}(\mathrm{MeV})$
due to insufficient
$\nu e \leftrightarrow \nu e$ & $e^- e^+ \leftrightarrow \nu\bar{\nu}$

Relic neutrinos in cosmology: $N_{\rm eff}$

Radiation energy density ρ_r in the early Universe:

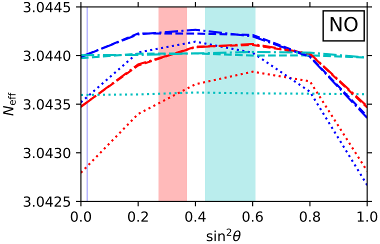
$$\rho_r = \left[1 + \frac{7}{8}\left(\frac{4}{11}\right)^{4/3} N_{\rm eff}\right]\rho_\gamma = [1 + 0.2271 N_{\rm eff}]\,\rho_\gamma$$

ρ_γ photon energy density, 7/8 is for fermions, $(4/11)^{4/3}$ due to photon reheating after neutrino decoupling

- $N_{\rm eff} \to$ all the radiation contribution not given by photons
- $N_{\rm eff} \simeq 1$ correspond to a single family of active neutrino, in equilibrium in the early Universe
- Active neutrinos:
  [Bennett, SG et al., 2020] [Froustey et al., 2020]: $N_{\rm eff} = 3.044$ See later!
  due to not instantaneous decoupling for the neutrinos
- $+$ Non Standard Interactions: $3.040 < N_{\rm eff} < 3.059$ [de Salas et al., 2016]

Observations: $N_{\rm eff} \simeq 3.0 \pm 0.2$ [Planck 2018]
Indirect probe of cosmic neutrino background!

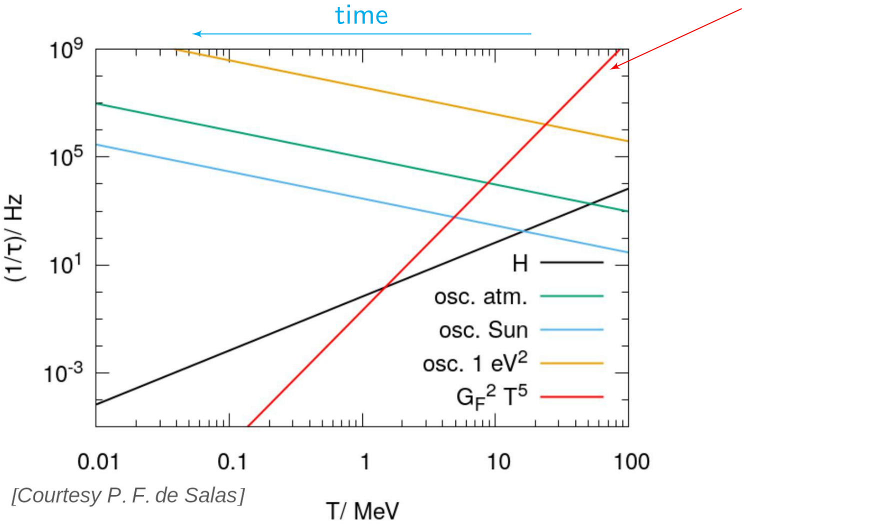
$\gg 10\sigma$!



$$N_{\rm eff} = 3.0440 \pm 0.0002$$

Neutrinos in the early Universe

before BBN: neutrinos coupled to plasma ($\nu_\alpha \bar{\nu}_\alpha \leftrightarrow e^+ e^-$, $\nu e \leftrightarrow \nu e$)



[Courtesy P. F. de Salas]

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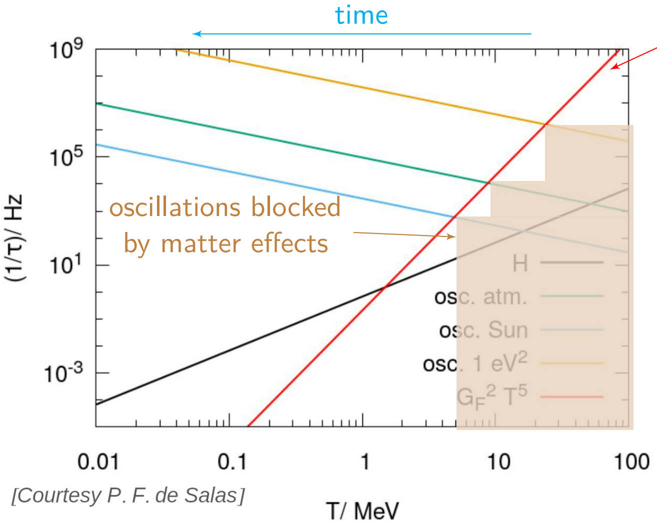

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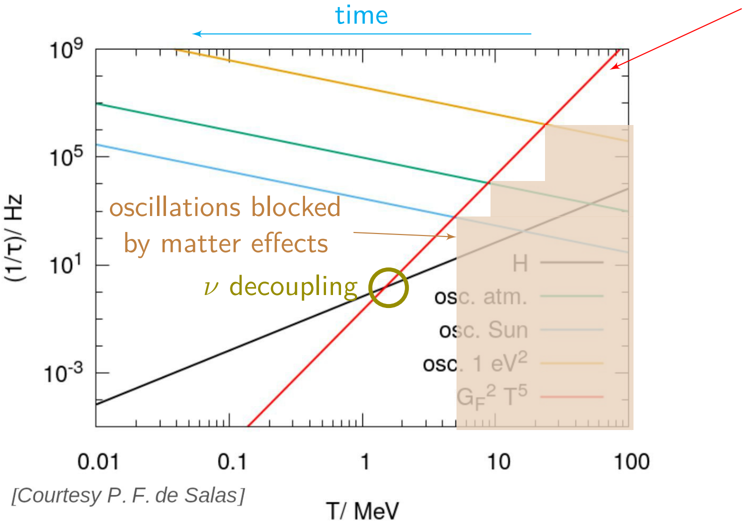

[Courtesy P. F. de Salas]

ν decouple mostly before $e^+e^- \to \gamma\gamma$ annihilation!

Neutrinos in the early Universe

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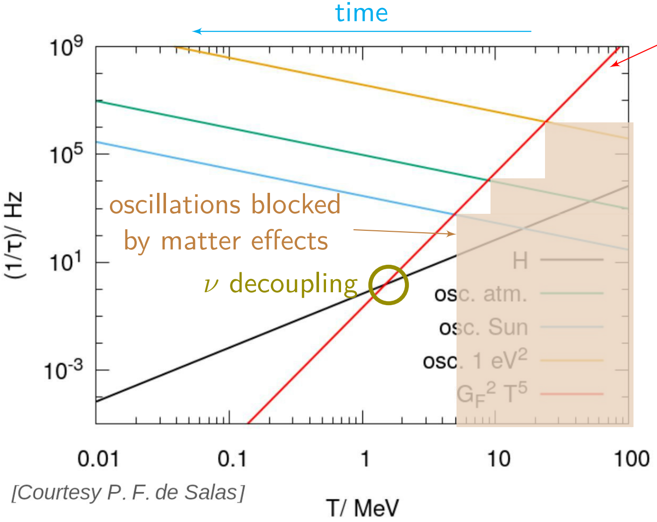

[Courtesy P. F. de Salas]

$T_\nu \simeq (4/11)^{1/3}\,T_\gamma$ after $e^+e^- \to \gamma\gamma$

f_ν: frozen Fermi-Dirac distribution

Today:

$T_{\nu,0} = 1.945$ K $\simeq 1.676 \times 10^{-4}$ eV

$\langle E_\nu \rangle \simeq 3.1\,T_{\nu,0} \simeq 5 \times 10^{-4}$ eV

$n_0 = n_{\nu,0} = n_{\bar\nu,0} \simeq 56$ cm^{-3} per family

ν decouple mostly before $e^+e^- \to \gamma\gamma$ annihilation!

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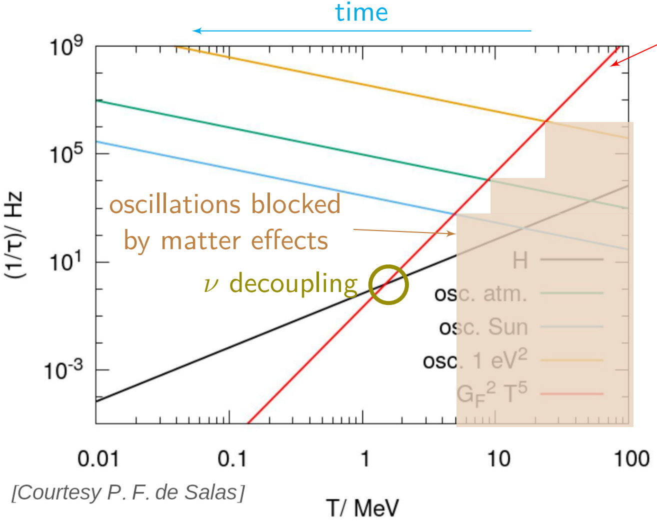

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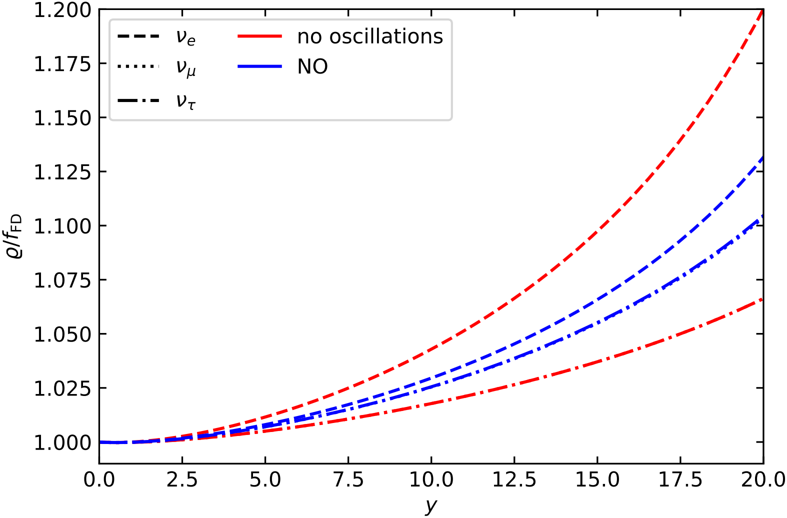
ν decouple mostly before $e^+e^- \to \gamma\gamma$ annihilation!

actually, the decoupling T is momentum dependent!

→ distortions to equilibrium f_ν!

Neutrino momentum distribution and $N_{\rm eff}$ [Bennett, SG+, 2012.02726]

Distortion of the momentum distribution ($f_{\rm FD}$: Fermi-Dirac at equilibrium)

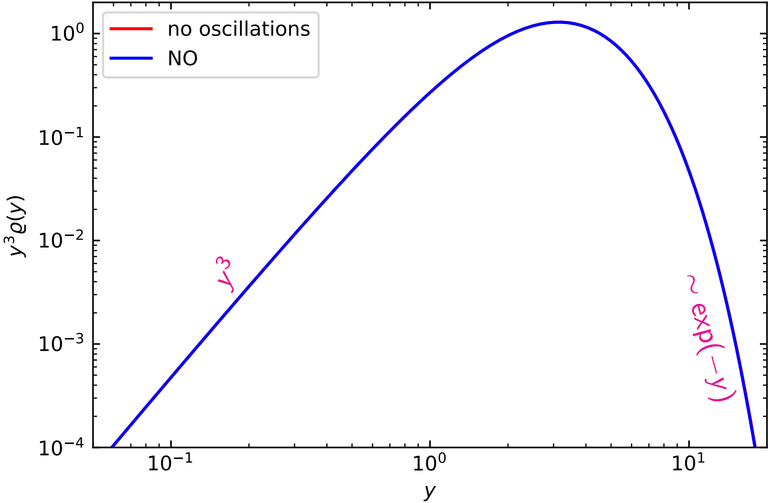


Neutrino momentum distribution and $N_{\rm eff}$ [Bennett, SG+, 2012.02726]

$$N_{\rm eff}^{\rm final} = \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\frac{\rho_\nu}{\rho_\gamma} = \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\frac{1}{\rho_\gamma}\sum_i g_i \int \frac{d^3p}{(2\pi)^3} E(p)\, f_{\nu,i}(p)$$

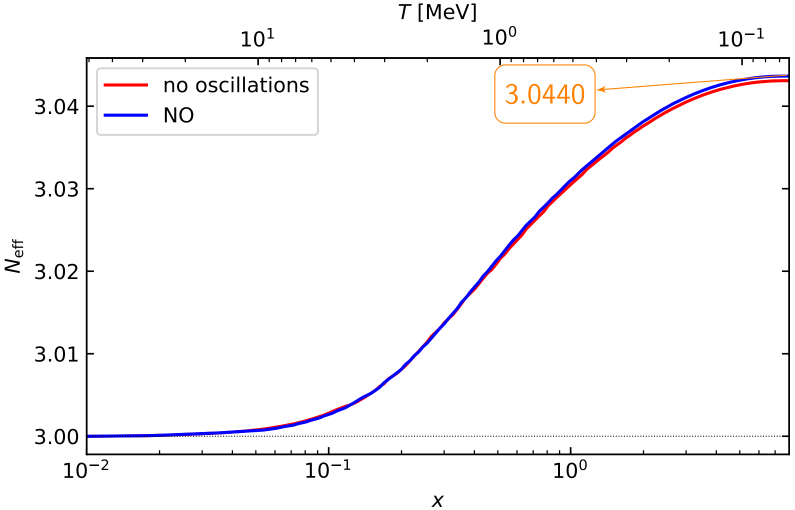
$(11/4)^{1/3} = (T_\gamma/T_\nu)^{\rm fin}$

$\hookrightarrow \propto y^3 \varrho_{ii}(y)$



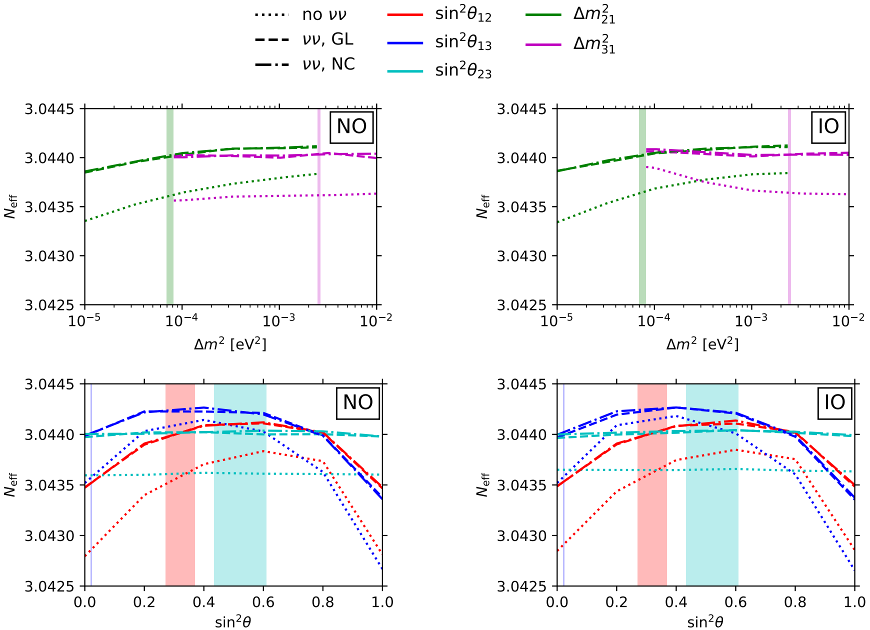
Neutrino momentum distribution and $N_{\rm eff}$ [Bennett, SG+, 2012.02726]

$$N_{\rm eff}^{\rm any\ time} = \frac{8}{7}\left(\frac{T_\gamma}{T_\nu}\right)^4 \frac{\rho_\nu}{\rho_\gamma} = \frac{8}{7}\left(\frac{T_\gamma}{T_\nu}\right)^4 \frac{1}{\rho_\gamma}\sum_i g_i \int \frac{d^3p}{(2\pi)^3} E(p)\, f_{\nu,i}(p)$$



Effect of neutrino oscillations

[Bennett, SG+, 2012.02726]



Effect of neutrino oscillations

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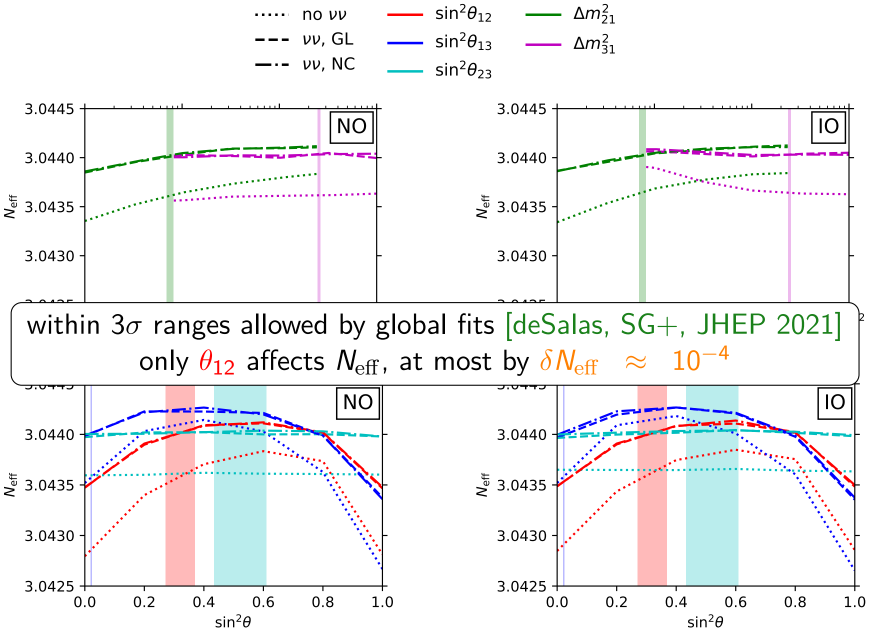

within 3σ ranges allowed by global fits [deSalas, SG+, JHEP 2021]
only θ_{12} affects $N_{\rm eff}$, at most by $\delta N_{\rm eff} \approx 10^{-4}$

Recommended $N_{\rm eff}$ value

[Bennett, SG+, 2012.02726]

Benchmark A: no $\nu\nu$ collisions

	$N_{\rm eff}^{\rm SM}$ (no osc)	$N_{\rm eff}^{\rm SM}$ (NO)	$N_{\rm eff}^{\rm SM}$ (IO)
Benchmark A — $\{\mathcal{I}_{\nu\nu}[\varrho]\}_{\alpha\alpha} = 0$			
Assuming: • (2)l̸n + (2) ln +(3)+ type (a) weak rates • Damping for $\{\mathcal{I}_{\nu e}[\varrho]\}_{\alpha\beta}$ • $N_y = 60$, $y_{\rm max} = 20$, NC linearly spaced y_i	**3.04263**	**3.04360**	**3.04361**
Alternative estimates			
Momentum grid			
$N_y = 40$, $y_{\rm max} = 20$, GL spacing of y_i nodes	3.04261	3.04355	3.04360
Integrals for off-diagonal $\{\mathcal{I}_{\nu e}[\varrho]\}_{\alpha\beta}$			
$N_y = 60$, $y_{\rm max} = 20$, NC linearly spaced y_i	3.04261	3.04357	3.04362
$N_y = 40$, $y_{\rm max} = 20$, GL spacing of y_i	3.04261	3.04357	3.04364
Finite-temperature QED corrections			
(2)l̸n	3.04361	3.04458	
(2)l̸n + (2) ln	3.04358	3.04452	
(2)l̸n + (3)	3.04264	3.04361	
(2)l̸n + (2) ln +(3)	3.04263	3.04360	

Recommended $N_{\rm eff}$ value

[Bennett, SG+, 2012.02726]

Benchmark B: full collision terms

	$N_{\rm eff}^{\rm SM}$ (no osc)	$N_{\rm eff}^{\rm SM}$ (NO)	$N_{\rm eff}^{\rm SM}$ (IO)
Benchmark B — $\{\mathcal{I}_{\nu\nu}[\varrho]\}_{\alpha\alpha} \neq 0$			
Assuming: • (2)$\not{\ln}$ + (2) ln +(3)+ type (a) weak rates • Full $\mathcal{I}_{\nu e}[\varrho]$ and $\mathcal{I}_{\nu\nu}[\varrho]$ • $N_y = 80$, $y_{\rm max} = 30$, NC linearly spaced y_i	**3.04341**	**3.04398**	**3.04399**
Alternative est			
Momentum grid			
$N_y = 80$, $y_{\rm max} = 30$, GL spacing of y_i	3.04334	3.04392	3.04392
$N_y = 80$, $y_{\rm max} = 20$, NC linearly spaced y_i	3.04334	3.04389	3.04391
$N_y = 80$, $y_{\rm max} = 20$, GL spacing of y_i	3.04334	3.04386	3.04393
Off-diagonal collision terms			
Damping terms, NC quadrature	3.04342	3.04408	
Damping terms, GL quadrature	3.04335	3.04399	
Neutrino–neutrino collision integral - $y_{\rm max} = 20$			
Diagonal ϱ	3.04333	3.04416	
Full ϱ, interpolate ϱ/FD only in diagonal	3.04334	3.04389	
Full ϱ, interpolate ϱ/FD also in off-diagonal	3.04334	3.04389	

$\nu\nu$ terms add $\sim (4 \div 8) \times 10^{-4}$

Recommended $N_{\rm eff}$ value

[Bennett, SG+, 2012.02726]

Benchmark B: full collision terms

	$N_{\rm eff}^{\rm SM}$ (no osc)	$N_{\rm eff}^{\rm SM}$ (NO)	$N_{\rm eff}^{\rm SM}$ (IO)
Benchmark B — $\{\mathcal{I}_{\nu\nu}[\varrho]\}_{\alpha\alpha} \neq 0$			
Assuming: • (2)~~ln~~ + (2) ln +(3)+ type (a) weak rates • Full $\mathcal{I}_{\nu e}[\varrho]$ and $\mathcal{I}_{\nu\nu}[\varrho]$ • $N_y = 80$, $y_{\rm max} = 30$, NC linearly spaced y_i	**3.04341**	**3.04398**	**3.04399**
Alternative estimates			
$N_y = 80$, $y_{\rm max}$			3.04392
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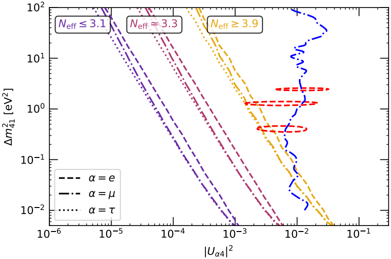
$N_y = 80$, $y_{\rm max}$			3.04393
Off-diagonal collision terms			
Damping t			
Damping t			
Neutrino–neutrino collision integral - $y_{\rm max} = 20$			
Diagonal ϱ	3.04333	3.04416	
Full ϱ, interpolate ϱ/FD only in diagonal	3.04334	3.04389	
Full ϱ, interpolate ϱ/FD also in off-diagonal	3.04334	3.04389	

Our recommended value (normal ordering):

$$N_{\rm eff} = 3.0440 \pm 0.0002$$

(numerical+physical uncertainty)

Full agreement with other results in literature
e.g. [Froustey+, JCAP 2020] & [Akita+, JCAP 2020]



$N_{\text{eff}} \leq 3.1$
$N_{\text{eff}} \simeq 3.3$
$N_{\text{eff}} \geq 3.9$
Δm^2_{41} [eV^2]
$|U_{\alpha 4}|^2$
$\alpha = e$
$\alpha = \mu$
$\alpha = \tau$

Short Baseline (SBL) anomaly

[SG+, JPG 43 (2016) 033001]

Problem: anomalies in SBL experiments → errors in flux calculations? / deviations from 3-ν description?

A short review:

LSND search for $\bar{\nu}_\mu \to \bar{\nu}_e$, with $L/E = 0.4 \div 1.5$ m/MeV. Observed a 3.8σ excess of $\bar{\nu}_e$ events [Aguilar et al., 2001]

Reactor re-evaluation of the expected anti-neutrino flux ⇒ disappearance of $\bar{\nu}_e$ events compared to predictions ($\sim 3\sigma$) with $L < 100$ m [Mention et al, 2011], [Azabajan et al, 2012]

Gallium calibration of GALLEX and SAGE Gallium solar neutrino experiments give a 2.7σ anomaly (disappearance of ν_e) [Giunti, Laveder, 2011]

MiniBooNE See next

Possible explanation:

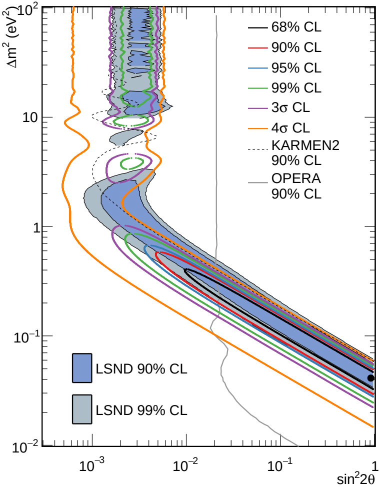
Additional squared mass difference $\Delta m^2_{\mathrm{SBL}} \simeq 1\ \mathrm{eV}^2$

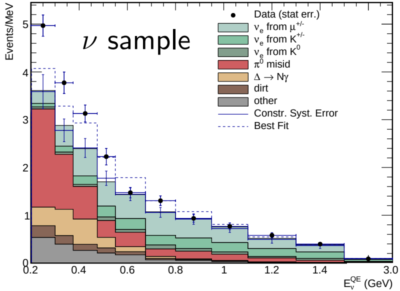
MiniBooNE

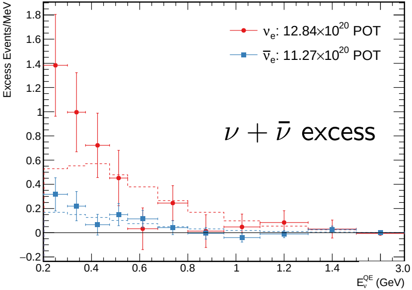
purpose: check LSND signal

$L \simeq 541$ m, 200 MeV $\leq E \lesssim 3$ GeV

no money, no near detector







MiniBooNE

[PRL 121 (2018) 221801]

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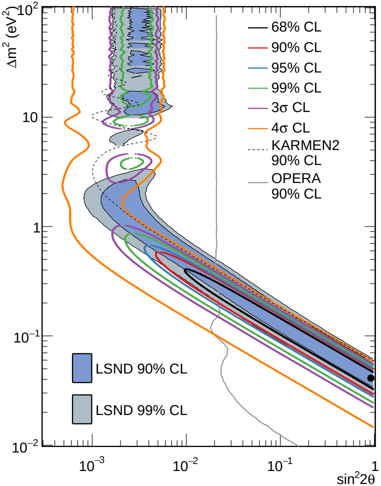
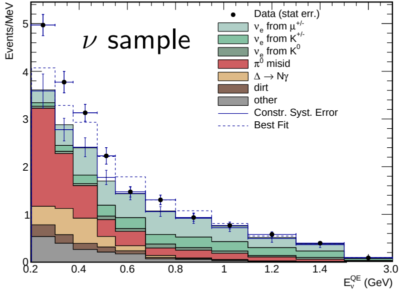
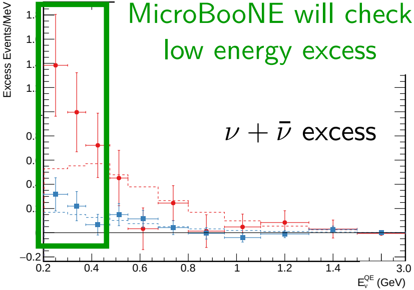


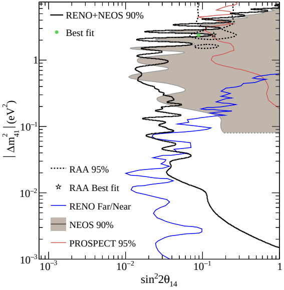




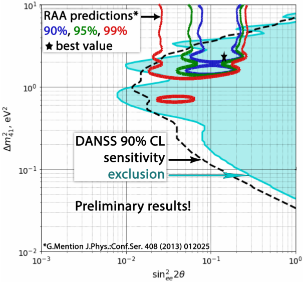
MicroBooNE will check low energy excess

ν_s at reactors in 2020

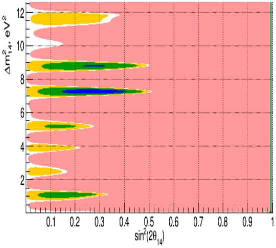
[RENO+NEOS, 2020]



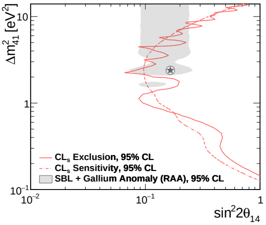
[DANSS, 2020]



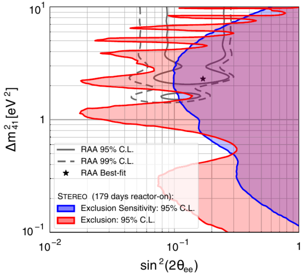
[Neutrino-4, PZETF 2020]



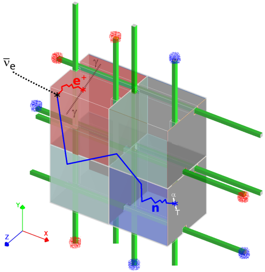
[PROSPECT, PRD 2020]



[STEREO, PRD 2020]



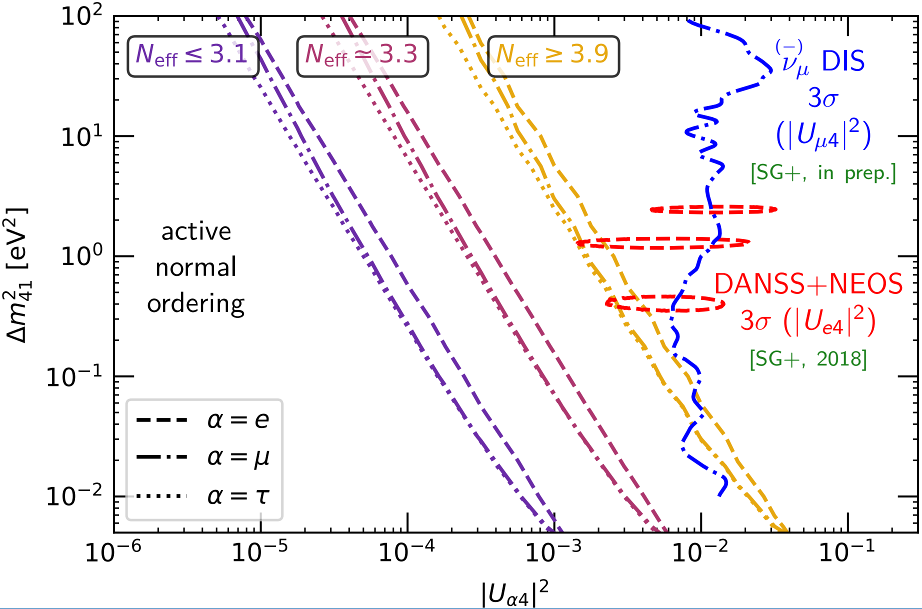
[SoLiD, JINST 2018]



$N_{\rm eff}$ and the new mixing parameters

[SG+, JCAP 07 (2019) 014]

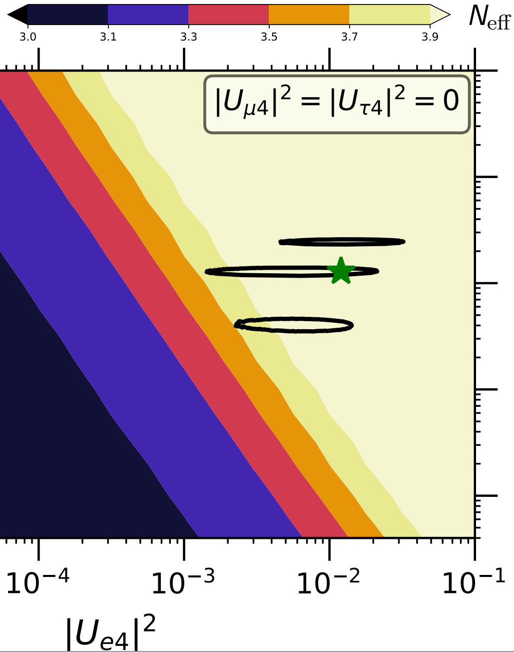
Only vary one angle and fix two to zero: do they have the same effect?



N_{eff} and the new mixing parameters

[SG+, JCAP 07 (2019) 014]

We can vary more than one angle:



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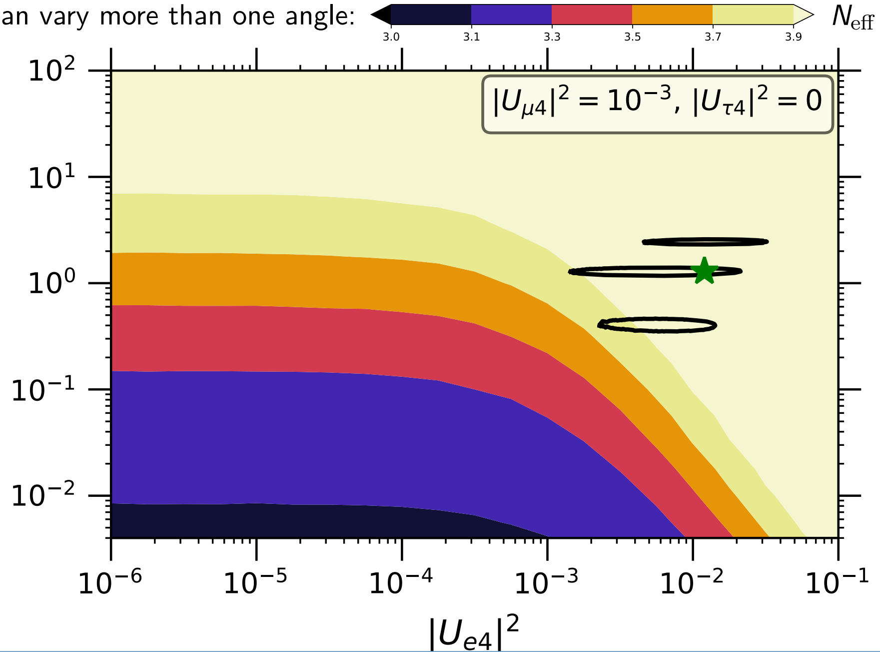

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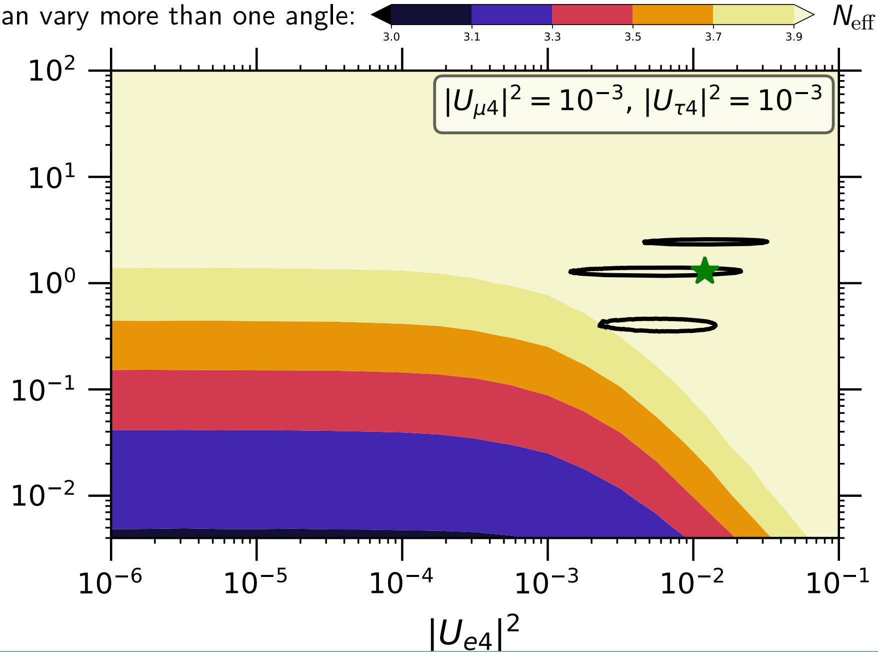
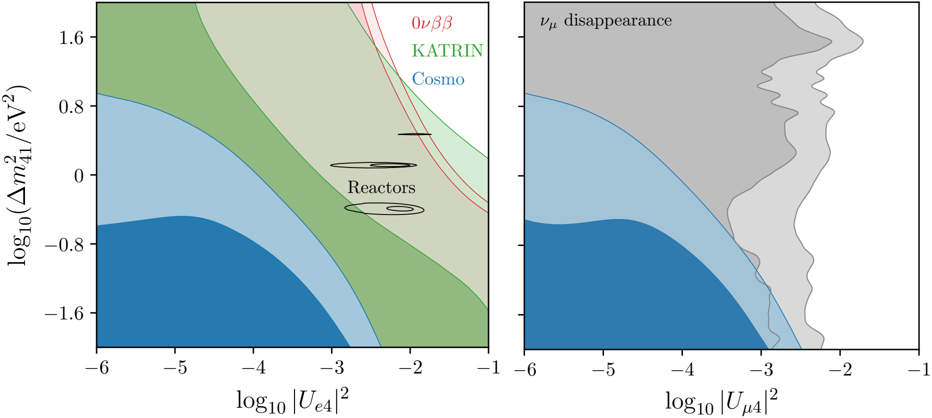


Comparing constraints

[arxiv:2003.02289]

Cosmological constraints are stronger than most other probes

But much more model dependent (as all the cosmological constraints)!



Comparing constraints

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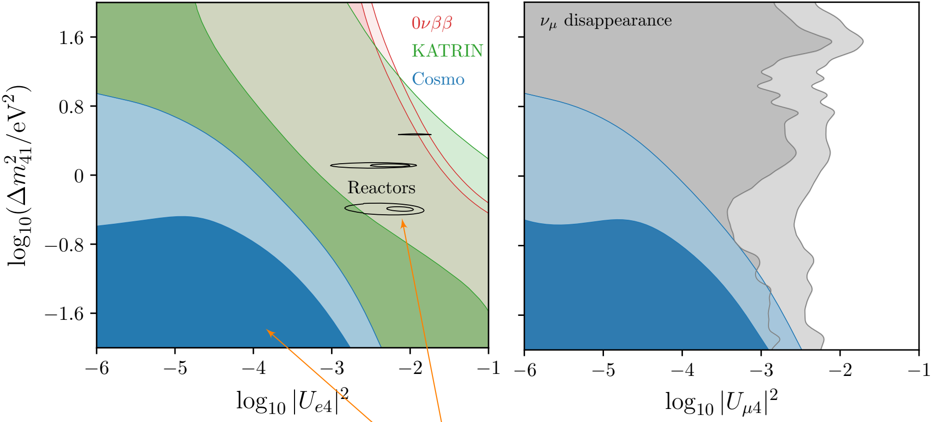
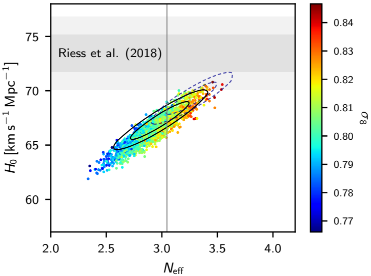
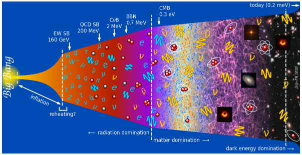


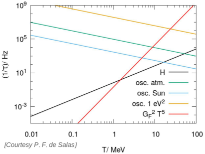
Warning: tension between reactor experiments and CMB bounds!

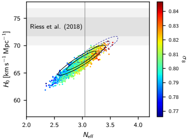

Riess et al. (2018)
H_0 [km s^{-1} Mpc^{-1}]
$N_{\rm eff}$
σ_8
75
70
65
60
2.0
2.5
3.0
3.5
4.0
0.84
0.83
0.82
0.81
0.80
0.79
0.78
0.77

Conclusions

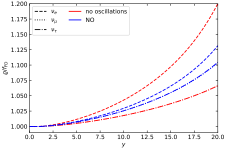
1 **Neutrinos in the early universe – probe lowest energies**

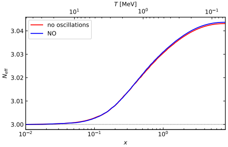


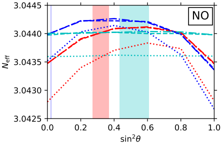




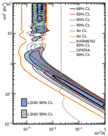
2 **Active neutrinos – precision**

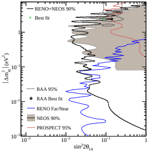


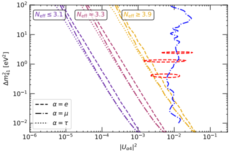


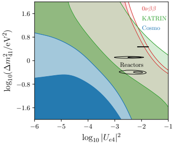


3 **Sterile neutrino hints – new physics?**



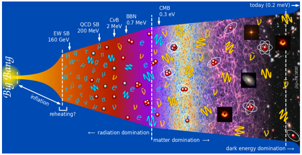


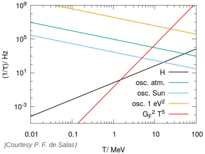


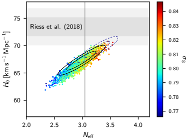


Conclusions

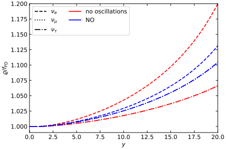
1 Neutrinos in the early universe – probe lowest energies

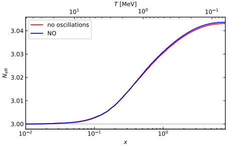


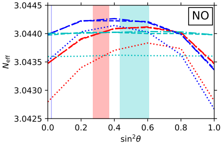




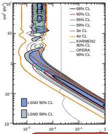
2 Active neutrinos – precision

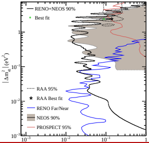


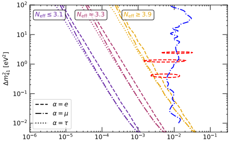


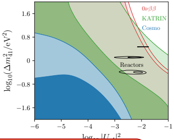


3 Sterile neutrino hints – new physics?









Thank you for the attention!

5 Backup

How precise is $N_{\rm eff} = 3.04\ldots$?

Long list of previous works...always less than 3ν mixing

[Mangano+, 2005]: $N_{\rm eff} = 3.046$	1st with 3ν mixing (still most cited value)
[de Salas+, 2016]: $N_{\rm eff} = 3.045$	updated collision terms
[SG+, 2019]: $N_{\rm eff} = 3.044$ FortEPiaNO code	more efficient and precise code, $N > 3$ neutrinos allowed, minor differences in numerical integrals
[Bennett+, 2019]: $N_{\rm eff} = 3.043$ (no full calculation)	finite-T QED corrections at $\mathcal{O}(e^3)$! further terms should be almost negligible
[Akita+, 2020]: $N_{\rm eff} = 3.044 \pm 0.0005$	equations in mass and flavor basis approximated $\nu\nu$ collisions
[Froustey+, 2020]: $N_{\rm eff} = 3.0440 \pm \mathcal{O}(10^{-4})$	full $\nu\nu$ interactions 1st estimate effect of CP-violating phase
[Bennett, SG+, 2020]: $N_{\rm eff} = 3.0440 \pm 0.0002$ FortEPiaNO improved	1st full discussion on effect of oscillation parameters, full estimation of current numerical and physical uncertainty

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

comoving coordinates: $a = 1/T \quad x \equiv m_e\, a \quad y \equiv p\, a \quad z \equiv T_\gamma\, a \quad w \equiv T_\nu\, a$

density matrix: $$\varrho(x,y) = \begin{pmatrix} \varrho_{ee} \equiv f_{\nu_e} & \varrho_{e\mu} & \varrho_{e\tau} \\ \varrho_{\mu e} & \varrho_{\mu\mu} \equiv f_{\nu_\mu} & \varrho_{\mu\tau} \\ \varrho_{\tau e} & \varrho_{\tau\mu} & \varrho_{\tau\tau} \equiv f_{\nu_\tau} \end{pmatrix}$$

$$\frac{\mathrm{d}\varrho(y,x)}{\mathrm{d}x} = \sqrt{\frac{3 m_{\rm Pl}^2}{8\pi\rho_{\rm T}}} \left\{ -i\frac{x^2}{m_e^3}\left[\frac{\mathbb{M}_{\rm F}}{2y} - \frac{2\sqrt{2}G_{\rm F}\, y}{x^6/m_e^6}\left(\frac{\mathbb{E}_\ell + \mathbb{P}_\ell}{m_W^2} + \frac{4\mathbb{E}_\nu}{3 m_Z^2} \right), \varrho \right] + \frac{m_e^3 G_{\rm F}^2}{(2\pi)^3 x^4 y^2}\mathcal{I}(\varrho) \right\}$$

$m_{\rm Pl}$ Planck mass – ρ_T total energy density – $m_{W,Z}$ mass of the W, Z bosons – $G_{\rm F}$ Fermi constant – $[.,.]$ commutator

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

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$$\frac{\mathrm{d}\varrho(y,x)}{\mathrm{d}x} = \sqrt{\frac{3m_{\mathrm{Pl}}^2}{8\pi\rho_{\mathrm{T}}}} \left\{ -i\frac{x^2}{m_e^3}\left[\frac{\mathbb{M}_{\mathrm{F}}}{2y} - \frac{2\sqrt{2}G_{\mathrm{F}}\, y}{x^6/m_e^6}\left(\frac{\mathbb{E}_\ell + \mathbb{P}_\ell}{m_W^2} + \frac{4\mathbb{E}_\nu}{3m_Z^2}\right), \varrho\right] + \frac{m_e^3 G_{\mathrm{F}}^2}{(2\pi)^3 x^4 y^2}\mathcal{I}(\varrho)\right\}$$

m_{Pl} Planck mass – ρ_T total energy density – $m_{W,Z}$ mass of the W, Z bosons – G_{F} Fermi constant – $[.,.]$ commutator

$$\mathbb{M}_{\mathrm{F}} = U \mathbb{M} U^\dagger$$

$$\mathbb{M} = \mathrm{diag}(m_1^2, \ldots, m_N^2)$$

$$U = R^{23} R^{13} R^{12}$$

e.g. $$R^{13} = \begin{pmatrix} \cos\theta_{13} & 0 & \sin\theta_{13} \\ 0 & 1 & 0 \\ -\sin\theta_{13} & 0 & \cos\theta_{13} \end{pmatrix}$$

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

comoving coordinates: $a = 1/T \quad x \equiv m_e\, a \quad y \equiv p\, a \quad z \equiv T_\gamma\, a \quad w \equiv T_\nu\, a$

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$$\mathbb{M}_{\rm F} = U \mathbb{M} U^\dagger$$

$$\mathbb{E}_\ell = \mathrm{diag}(\rho_e, \rho_\mu, 0) \qquad \mathbb{E}_\nu = S_a\left(\int dy y^3 \varrho\right) S_a \quad \text{with } S_a = \mathrm{diag}(1,1,1)$$

lepton densities — neutrino densities — (only for active neutrinos)

take into account matter effects in oscillations

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

comoving coordinates: $a = 1/T$ $\quad x \equiv m_e\, a$ $\quad y \equiv p\, a$ $\quad z \equiv T_\gamma\, a$ $\quad w \equiv T_\nu\, a$

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m_{Pl} Planck mass – ρ_T total energy density – $m_{W,Z}$ mass of the W, Z bosons – G_{F} Fermi constant – $[.,.]$ commutator

$\mathbb{M}_{\mathrm{F}} = U \mathbb{M} U^\dagger$ $\qquad \mathbb{E}_\ell = \mathrm{diag}(\rho_e, \rho_\mu, 0)$ $\qquad \mathbb{E}_\nu = S_a \left(\int dy y^3 \varrho \right) S_a$

$\mathcal{I}(\varrho)$ collision integrals

take into account neutrino-electron scattering and pair annihilation, plus neutrino–neutrino interactions

2D integrals over momentum, take most of the computation time

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

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$\mathcal{I}(\varrho)$ collision integrals

from continuity equation
$\dot\rho = -3H(\rho + P)$

$$\frac{\mathrm{d}z}{\mathrm{d}x} = \frac{\sum_{\ell=e,\mu}\left[\frac{r_\ell^2}{r}J(r_\ell)\right] + G_1(r) - \frac{1}{2\pi^2 z^3}\int_0^\infty dy\, y^3 \sum_{\alpha=e}^{\tau}\frac{\mathrm{d}\varrho_{\alpha\alpha}}{\mathrm{d}x}}{\sum_{\ell=e,\mu}\left[r_\ell^2 J(r_\ell) + Y(r_\ell)\right] + G_2(r) + \frac{2\pi^2}{15}}$$

$r = x/z$, $r_\ell = m_\ell/m_e\, r$ $\qquad$ $J(r)$, $Y(r)$ from non-relativistic transition of $e^\pm$, $\mu^\pm$
$G_1(r)$ and $G_2(r)$ from electromagnetic corrections

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

comoving coordinates: $a = 1/T \quad x \equiv m_e\, a \quad y \equiv p\, a \quad z \equiv T_\gamma\, a \quad w \equiv T_\nu\, a$

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neutrino temperature w: same equation as z, but electrons always relativistic

ν oscillations in the early universe

[Bennett, SG+, 2012.02726]

comoving coordinates: $a = 1/T \quad x \equiv m_e\, a \quad y \equiv p\, a \quad z \equiv T_\gamma\, a \quad w \equiv T_\nu\, a$

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$$\frac{\mathrm{d}\varrho(y,x)}{\mathrm{d}x} = \sqrt{\frac{3m_{\mathrm{Pl}}^2}{8\pi\rho_{\mathrm{T}}}} \left\{ -i\frac{x^2}{m_e^3}\left[\frac{\mathbb{M}_{\mathrm{F}}}{2y} - \frac{2\sqrt{2}G_{\mathrm{F}}\,y}{x^6/m_e^6}\left(\frac{\mathbb{E}_\ell + \mathbb{P}_\ell}{m_W^2} + \frac{4\mathbb{E}_\nu}{3m_Z^2}\right), \varrho\right] + \frac{m_e^3 G_{\mathrm{F}}^2}{(2\pi)^3 x^4 y^2}\mathcal{I}(\varrho)\right\}$$

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$$\mathbb{M}_{\mathrm{F}} = U\mathbb{M}U^\dagger \qquad \mathbb{E}_\ell = \mathrm{diag}(\rho_e, \rho_\mu, 0) \qquad \mathbb{E}_\nu = S_a\left(\int dy y^3 \varrho\right) S_a$$

$\mathcal{I}(\varrho)$ collision integrals

from continuity equation $\dot{\rho} = -3H(\rho + P)$

$$\frac{\mathrm{d}z}{\mathrm{d}x} = \frac{\sum_{\ell=e,\mu}\left[\frac{r_\ell^2}{r}J(r_\ell)\right] + G_1(r) - \frac{1}{2\pi^2 z^3}\int_0^\infty dy\, y^3 \sum_{\alpha=e}^{\tau}\frac{\mathrm{d}\varrho_{\alpha\alpha}}{\mathrm{d}x}}{\sum_{\ell=e,\mu}\left[r_\ell^2 J(r_\ell) + Y(r_\ell)\right] + G_2(r) + \frac{2\pi^2}{15}}$$

neutrino temperature w: same equation as z, but electrons always relativistic

initial conditions: $\varrho_{\alpha\alpha}$ = Fermi-Dirac at $x_{\mathrm{in}} \simeq 0.001$, with $w = z \simeq 1$

ν oscillations in the early universe

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m_{Pl}: Planck mass, ρ_{T}: total energy density, $m_{W,Z}$: mass of the W, Z bosons, G_{F}: Fermi constant, $[\,,\,]$ commutator

FORTran-Evolved PrimordIAl Neutrino Oscillations (FortEPiaNO)
https://bitbucket.org/ahep_cosmo/fortepiano_public

from continuity equation
$\dot\rho = -3H(\rho + P)$

$$\frac{\mathrm{d}z}{\mathrm{d}x} = \frac{\ldots - \frac{1}{2\pi^2 z^3}\int_0 dy\, y^3 \sum_{\alpha=e} \frac{\mathrm{d}\varrho_{\alpha\alpha}}{\mathrm{d}x}}{\sum_{\ell=e,\mu}\left[r_\ell^2 J(r_\ell) + Y(r_\ell)\right] + G_2(r) + \frac{2\pi^2}{15}}$$

will be public soon

neutrino temperature w: same equation as z, but electrons always relativistic

initial conditions: $\varrho_{\alpha\alpha}$ = Fermi-Dirac at $x_{\mathrm{in}} \simeq 0.001$, with $w = z \simeq 1$

Neutrino-neutrino interactions

[Bennett, SG+, 2012.02726]

Contribution to collision terms:

$$\mathcal{I}_{\nu\nu}[\varrho(y)] \propto G_F^2 \int \mathrm{d}y_2\, \mathrm{d}y_3\, \Pi_{\nu\nu}(y, y_2, y_3; x)\, F_{\nu\nu}(\varrho(y), \varrho(y_2), \varrho(y_3), \varrho(y_4))$$

$\Pi_{\nu\nu}(y, y_2, y_3; x)$: integrals of some combination of neutrino momenta

Critical function: $F_{\nu\nu}$!

it contains combinations such as $\varrho^{(1)}\varrho^{(3)}\varrho^{(2)}\varrho^{(4)}$ and permutations

it increases complexity of the code!

couples modes non-linearly

numerically more expensive (stronger dependence on y_i grid than νe terms)

Neutrino-neutrino interactions

[Bennett, SG+, 2012.02726]

Contribution to collision terms:

$$\mathcal{I}_{\nu\nu}[\varrho(y)] \propto G_F^2 \int \mathrm{d}y_2\, \mathrm{d}y_3\, \Pi_{\nu\nu}(y, y_2, y_3; x)\, F_{\nu\nu}(\varrho(y), \varrho(y_2), \varrho(y_3), \varrho(y_4))$$

	$N_{\rm eff}^{\rm SM}$ (no osc)	$N_{\rm eff}^{\rm SM}$ (NO)
Benchmark A — $\{\mathcal{I}_{\nu\nu}[\varrho]\}_{\alpha\alpha} = 0$		
Assuming: • $(2)\not\ln + (2)\ln + (3)+$ type (a) weak rates • Damping for $\{\mathcal{I}_{\nu e}[\varrho]\}_{\alpha\beta}$ • $N_y = 60$, $y_{\rm max} = 20$, NC linearly spaced y_i	**3.04263**	**3.04360**
— $\{\mathcal{I}_{\nu\nu}[\varrho]\}_{\alpha\alpha} \neq 0$		
Assuming: • $(2)\not\ln + (2)\ln + (3)+$ type (a) weak rates • Full $\mathcal{I}_{\nu e}[\varrho]$ and $\mathcal{I}_{\nu\nu}[\varrho]$ • $N_y = 80$, $y_{\rm max} = 30$, NC linearly spaced y_i	**3.04341**	**3.04398**
Neutrino–neutrino collision integral - $y_{\rm max} = 20$		
Diagonal ϱ	3.04333	3.04416
Full ϱ, interpolate ϱ/FD only in diagonal	3.04334	3.04389
Full ϱ, interpolate ϱ/FD also in off-diagonal	3.04334	3.04389

$\mathcal{I}_{\nu\nu}[\varrho(y)]$ is important! $(4 \div 8) \times 10^{-4}$

approximations may work

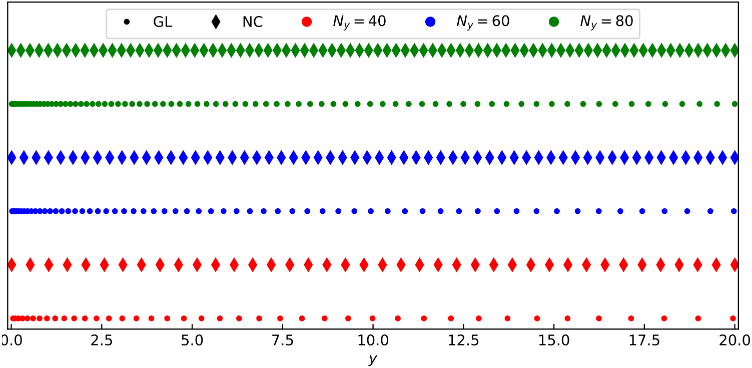
Sampling the y momenta

Discretize neutrino momenta to compute integrals and evolution

two sampling methods for y_i, with $i = 1, \ldots, N_y$:

linear spacing,
Newton-Cotes (NC) integration

Gauss-Laguerre (GL)
optimized for computing $\int_0^\infty dy\, f(y) e^{-y}$



Need to define range ($y_{\rm min} \le y \le y_{\rm max}$) and number of nodes N_y

Sampling the y momenta

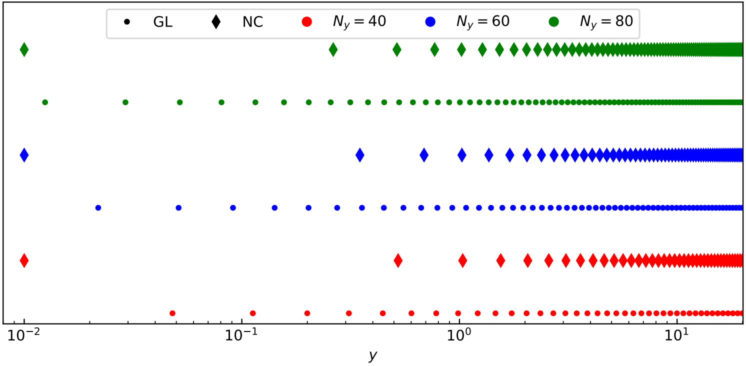
[Bennett, SG+, 2012.02726]

Discretize neutrino momenta to compute integrals and evolution

two sampling methods for y_i, with $i = 1, \ldots, N_y$:

linear spacing,
Newton-Cotes (NC) integration

Gauss-Laguerre (GL)
optimized for computing $\int_0^\infty dy\, f(y) e^{-y}$



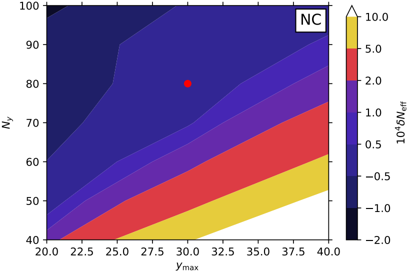
Need to define range ($y_{\min} \leq y \leq y_{\max}$) and number of nodes N_y

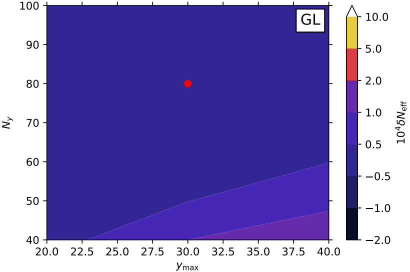
Sampling the y momenta

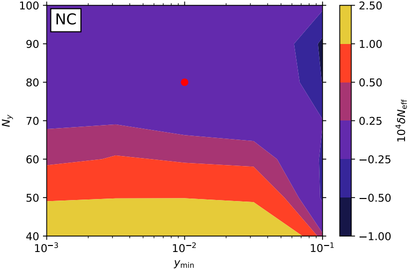
[Bennett, SG+, 2012.02726]

Discretize neutrino momenta to compute integrals and evolution

Results may depend on $y_{\rm min}$, $y_{\rm max}$, N_y





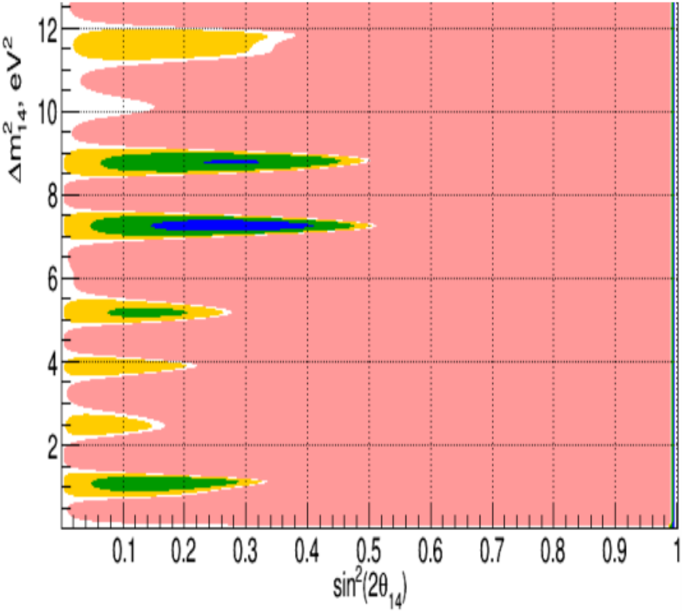


at same N_y,
GL results are more stable!

GL is more efficient

$\delta N_{\rm eff} \approx 10^{-4}$ from varying N_y, $y_{\rm max}$

Neutrino-4

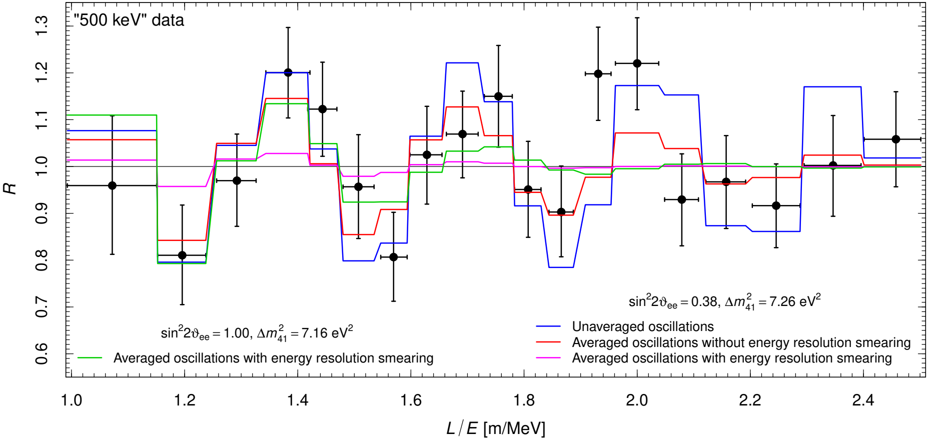


claimed $> 3\sigma$ preference for 3+1 over 3ν case

best fit incompatible with other reactor experiments

Neutrino-4

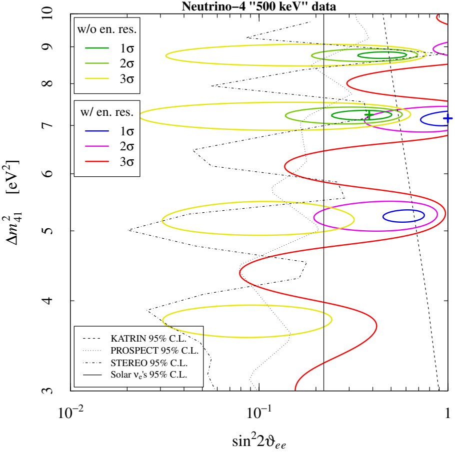
[Giunti+, arxiv:2101.06785]

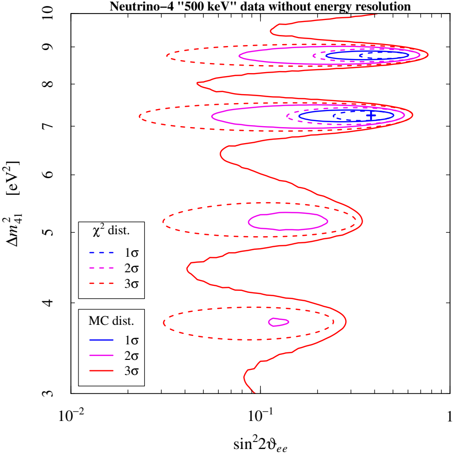


energy resolution smearing not properly taken into account?

Neutrino-4

[Giunti+, arxiv:2101.06785]





proper energy resolution treatment moves best-fit $\rightarrow \sin^2 2\vartheta \simeq 1$

need to take into account
violation of Wilk's theorem
↓
relaxed constraints