"la Caixa" Foundation
Junior Leader
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New neutrino physics with early universe probes

(Posthumous) FELLINI Seminar, 19/10/2023

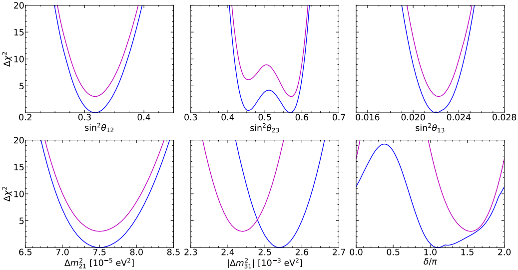
A

Active neutrinos

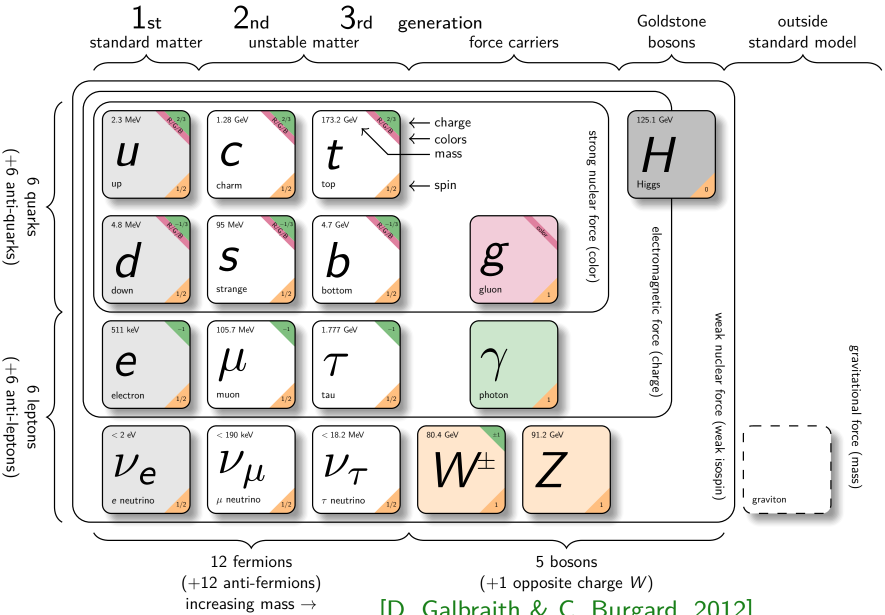
Spoiler: “Sterile” will come later

Based on:

- JHEP 02 (2021) 071 and update
- Planck 2018
- JCAP 04 (2021) 073

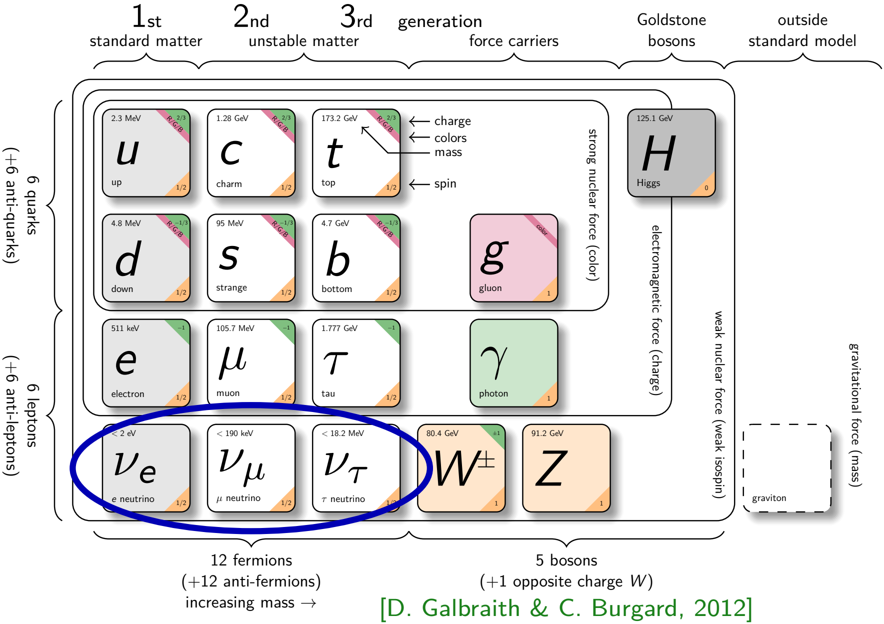


The Standard Model of Particle Physics



[D. Galbraith & C. Burgard, 2012]

The Standard Model of Particle Physics



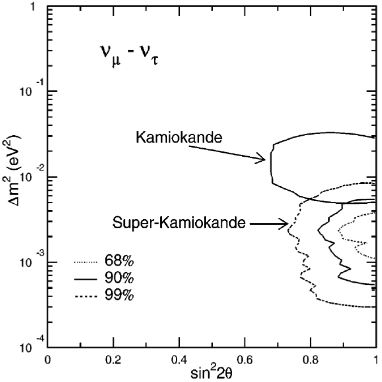
[D. Galbraith & C. Burgard, 2012]

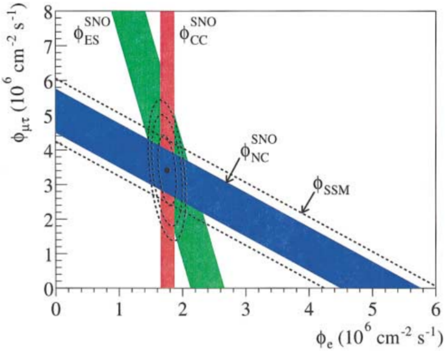
Neutrino oscillations

Major discoveries:

[SuperKamiokande, 1998]

[SNO, 2001-2002]





first discovery of $\nu_\mu \rightarrow \nu_\tau$ oscillations from atmospheric ν

first discovery of $\nu_e \rightarrow \nu_\mu, \nu_\tau$ oscillations from solar ν

Nobel prize in 2015

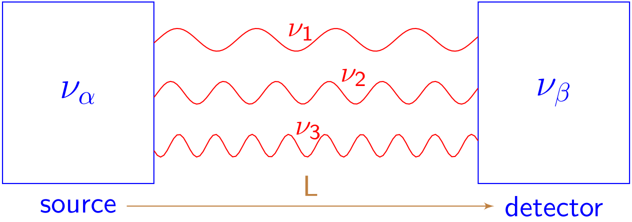
Two neutrino bases

flavor neutrinos ν_α

$|\nu_\alpha\rangle = \sum_k U_{\alpha k}|\nu_k\rangle$

massive neutrinos ν_k

$$|\nu(t=0)\rangle = |\nu_\alpha\rangle = U_{\alpha 1}|\nu_1\rangle + U_{\alpha 2}|\nu_2\rangle + U_{\alpha 3}|\nu_3\rangle$$



$$|\nu(t>0)\rangle = |\nu_\beta\rangle = U_{\alpha 1}\, e^{-iE_1 t}|\nu_1\rangle + U_{\alpha 2}\, e^{-iE_2 t}|\nu_2\rangle + U_{\alpha 3}\, e^{-iE_3 t}|\nu_3\rangle \neq |\nu_\alpha\rangle$$

$E_k^2 = p^2 + m_k^2 \longleftarrow$ define $\longrightarrow t = L$

$$\boxed{P_{\nu_\alpha\to\nu_\beta}(L) = |\langle\nu_\alpha|\nu(L)\rangle|^2 = \sum_{k,j} U_{\beta k} U^*_{\alpha k} U^*_{\beta j} U_{\alpha j} \exp\left(-i\frac{\Delta m^2_{kj} L}{2E}\right)}$$

$$\Delta m^2_{ij} = m^2_i - m^2_j$$

The mixing matrix

U can be parameterized using 3 angles (θ_{12}, θ_{13}, θ_{23}) and max 3 (1 Dirac δ, 2 Majorana [∃ only for Majorana ν]) phases

$$U = \underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\text{mainly atmospheric and LBL accelerator disappearance}} \underbrace{\begin{pmatrix} c_{13} & 0 & s_{13}e^{-i\delta} \\ 0 & 1 & 0 \\ -s_{13}e^{i\delta} & 0 & c_{13} \end{pmatrix}}_{\text{mainly LBL reactors and LBL accelerator appearance}} \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{mainly solar and VLBL reactors}} M$$

Majorana phases irrelevant for oscillation experiments

Relevant for example in neutrinoless double-beta decay

$$s_{ij} \equiv \sin\theta_{ij}; \; c_{ij} \equiv \cos\theta_{ij}$$

LBL = long baseline; VLBL = very long baseline;

Three Neutrino Oscillations

$$\nu_\alpha = \sum_{k=1}^{3} U_{\alpha k} \nu_k \quad (\alpha = e, \mu, \tau)$$

$U_{\alpha k}$ described by 3 mixing angles θ_{12}, θ_{13}, θ_{23} and one CP phase δ

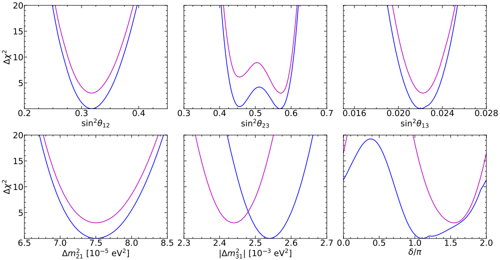
Current knowledge of the 3 active ν mixing: [JHEP 02 (2021) update]

NO/NH: Normal Ordering/Hierarchy, $m_1 < m_2 < m_3$ IO/IH: Inverted O/H, $m_3 < m_1 < m_2$

$$\begin{aligned}
\Delta m^2_{21} &= (7.50^{+0.22}_{-0.20}) \cdot 10^{-5}\ \mathrm{eV}^2 \\
|\Delta m^2_{31}| &= (2.54 \pm 0.03) \cdot 10^{-3}\ \mathrm{eV}^2\ (\mathrm{NO}) \\
&= (2.44 \pm 0.03) \cdot 10^{-3}\ \mathrm{eV}^2\ (\mathrm{IO})
\end{aligned}$$

$$\begin{aligned}
10 \sin^2(\theta_{12}) &= 3.18 \pm 0.16 \\
10^2 \sin^2(\theta_{13}) &= 2.200^{+0.069}_{-0.062}\ (\mathrm{NO}) \\
&= 2.225^{+0.064}_{-0.070}\ (\mathrm{IO}) \\
10 \sin^2(\theta_{23}) &= 4.55 \pm 0.13 U 5.71 \pm 0.12\ (\mathrm{NO}) \\
&= 5.71^{+0.14}_{-0.17}\ (\mathrm{IO})
\end{aligned}$$

$$\begin{aligned}
\delta/\pi &= 1.10^{+0.27}_{-0.12}\ (\mathrm{NO}) \\
&= 1.54 \pm 0.14\ (\mathrm{IO})
\end{aligned}$$

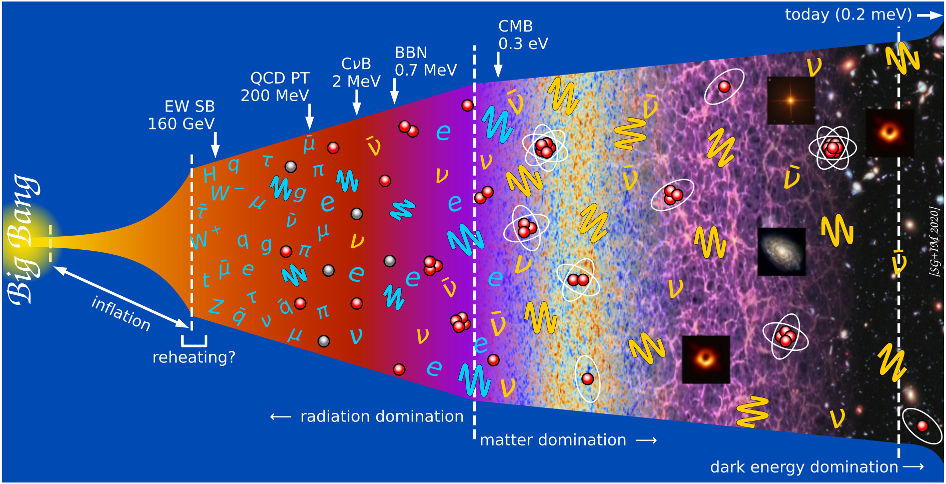


mass ordering
still unknown

δ still unknown

see also: `http://globalfit.astroparticles.es`

History of the universe



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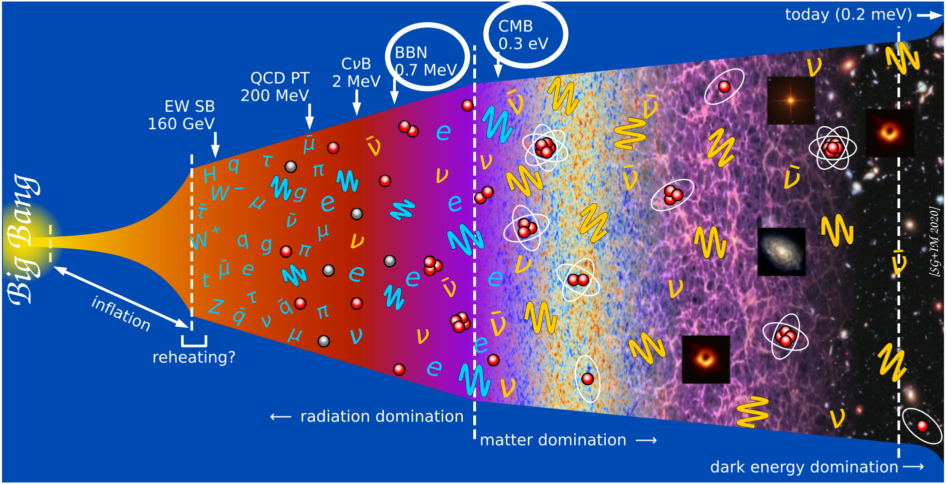

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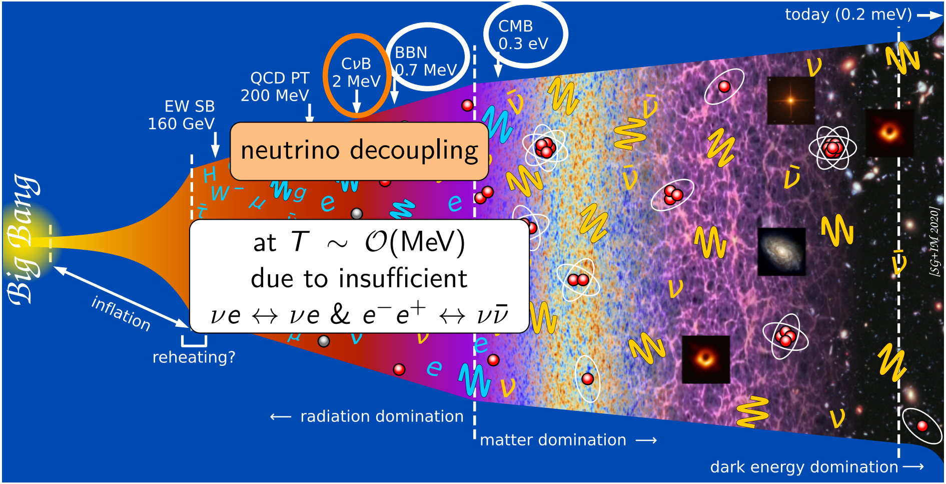


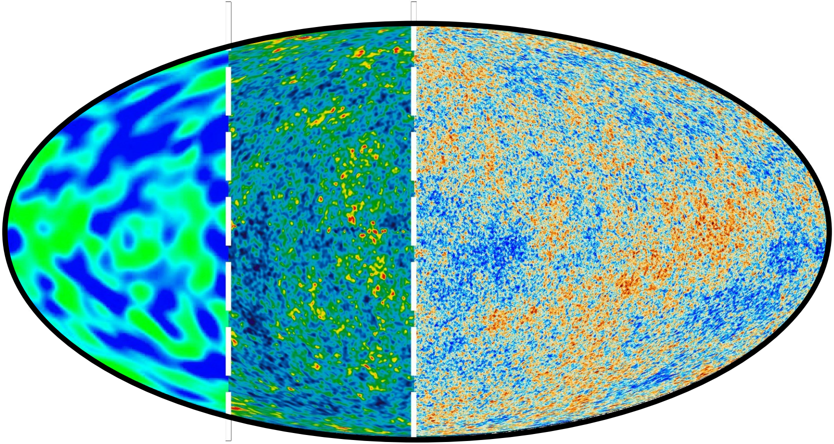
neutrino decoupling

at $T \sim \mathcal{O}(\text{MeV})$
due to insufficient
$\nu e \leftrightarrow \nu e$ & $e^- e^+ \leftrightarrow \nu\bar{\nu}$

The oldest picture of the Universe

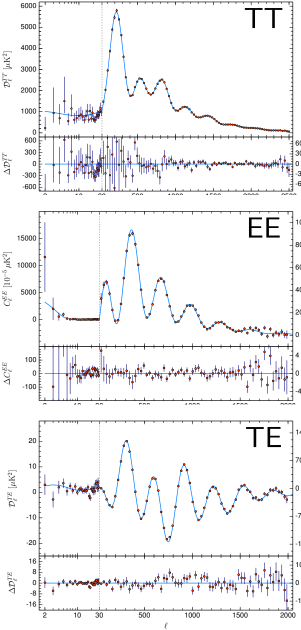
The Cosmic Microwave Background, generated at $t \simeq 4 \times 10^5$ years

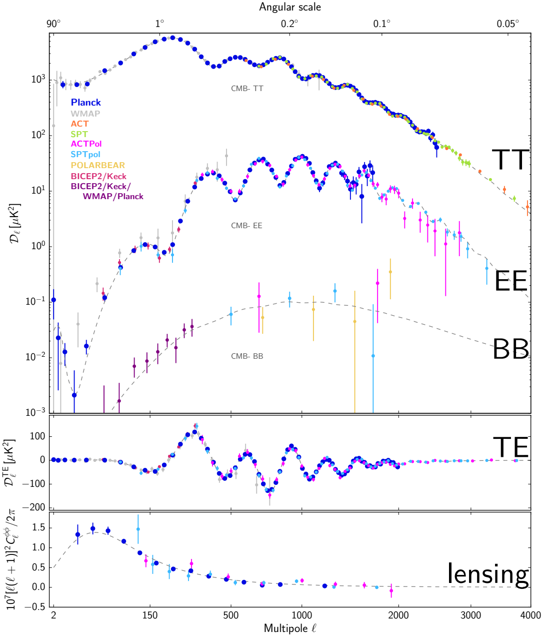
COBE (1992) WMAP (2003) Planck (2013)



CMB spectra as of 2018

[Planck Collaboration, 2018]





Big Bang Nucleosynthesis (BBN)

BBN: production of light nuclei at $t \sim 1$s to $t \sim \mathcal{O}(10^2)$s

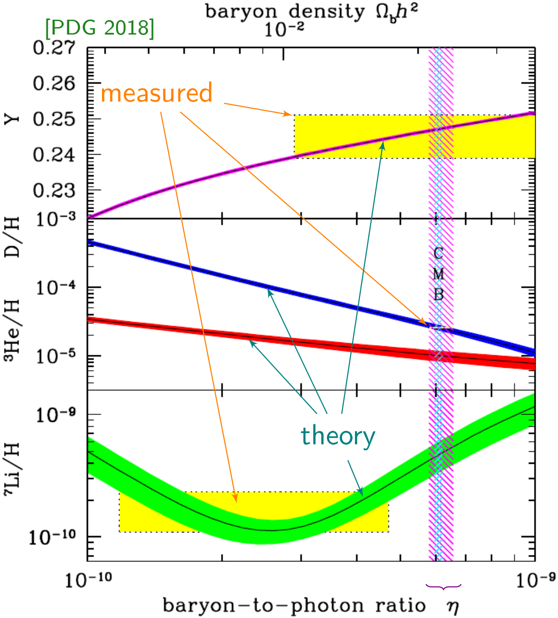
temperature $T_{fr} \simeq 1$ MeV
from nucleon freeze-out

much earlier than CMB!

strong probe for physics
before the CMB

e.g. neutrinos!

ν affect
universe expansion
and
reaction rates $(\nu_e/\bar{\nu}_e)$
at BBN time...



BBN concordance

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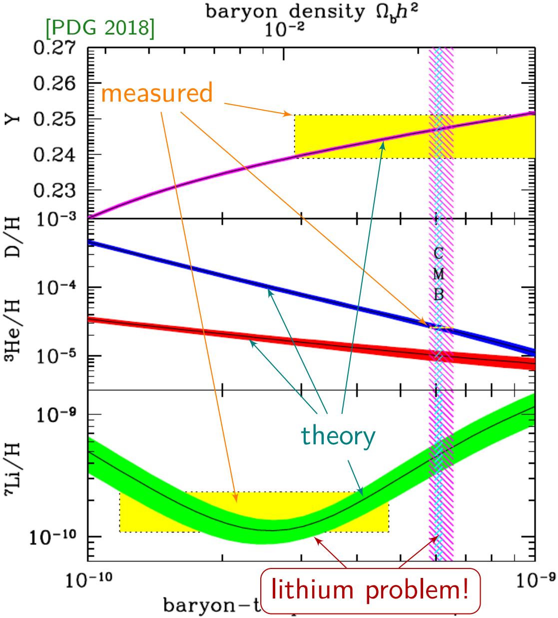


[PDG 2018]

measured

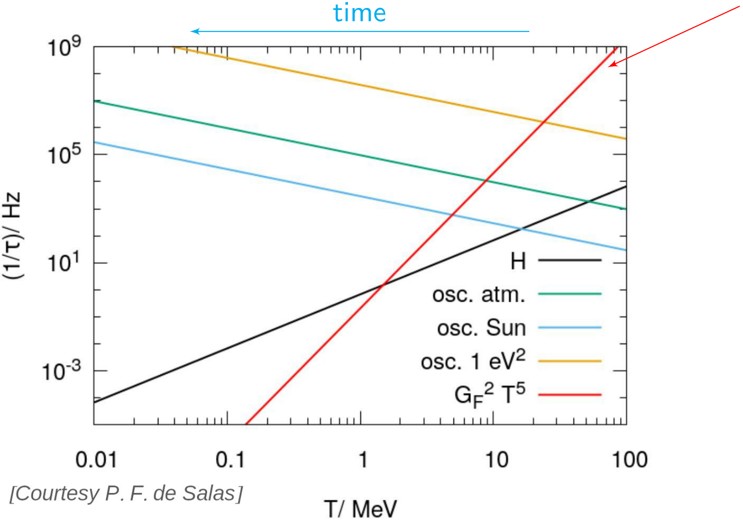
theory

lithium problem!

BBN concordance

Neutrinos in the early Universe

before BBN: neutrinos coupled to plasma ($\nu_\alpha\bar{\nu}_\alpha \leftrightarrow e^+e^-$, $\nu e \leftrightarrow \nu e$)



[Courtesy P. F. de Salas]

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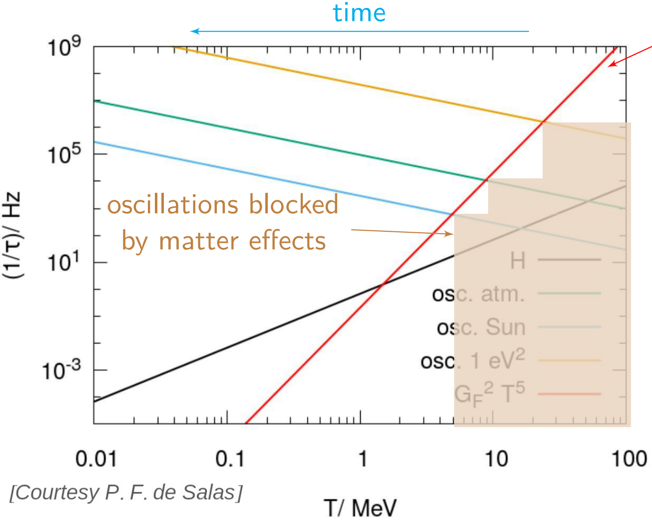

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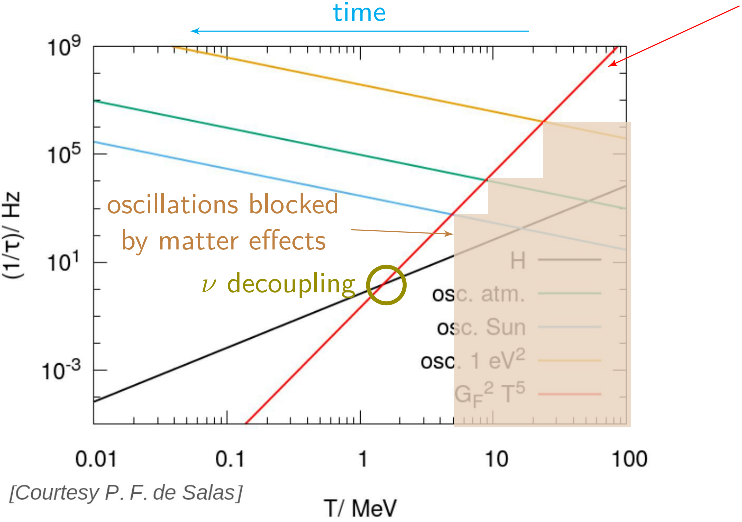

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ν decouple mostly before $e^+e^- \to \gamma\gamma$ annihilation!

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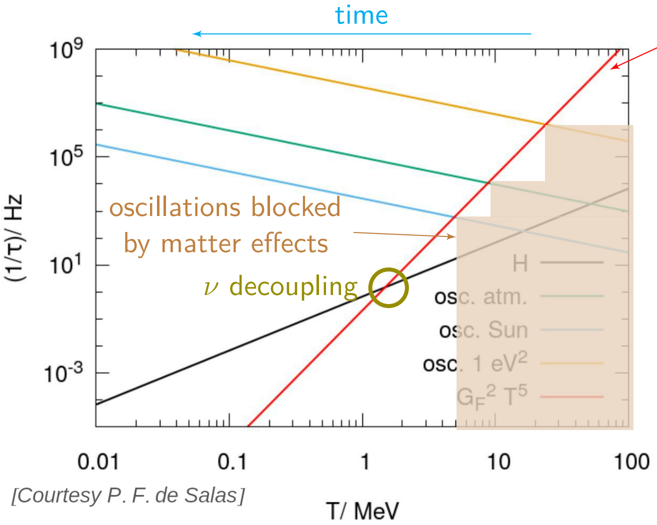

[Courtesy P. F. de Salas]

$T_\nu \simeq (4/11)^{1/3} T_\gamma$ after $e^+ e^- \to \gamma\gamma$

f_ν: frozen Fermi-Dirac distribution

Today:

$T_{\nu,0} = 1.945$ K $\simeq 1.676 \times 10^{-4}$ eV

$\langle E_\nu \rangle \simeq 3.1 T_{\nu,0} \simeq 5 \times 10^{-4}$ eV

$n_0 = n_{\nu,0} = n_{\bar{\nu},0} \simeq 56$ cm^{-3} per family

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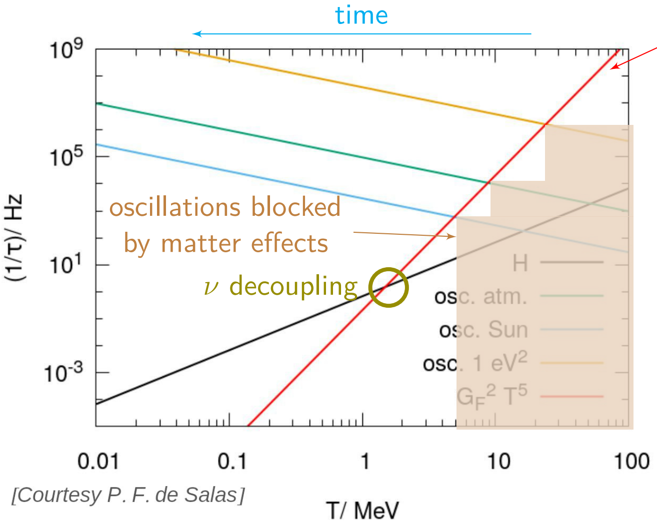

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actually, the decoupling T is momentum dependent! →

distortions to equilibrium f_ν!

ν oscillations in the early universe

[Bennett, SG+, JCAP 2021]
[Sigl, Raffelt, 1993]

comoving coordinates: $a = 1/T \quad x \equiv m_e\, a \quad y \equiv p\, a \quad z \equiv T_\gamma\, a \quad w \equiv T_\nu\, a$

density matrix: $\varrho(x,y) = \begin{pmatrix} \varrho_{ee} \equiv f_{\nu_e} & \varrho_{e\mu} & \varrho_{e\tau} \\ \varrho_{\mu e} & \varrho_{\mu\mu} \equiv f_{\nu_\mu} & \varrho_{\mu\tau} \\ \varrho_{\tau e} & \varrho_{\tau\mu} & \varrho_{\tau\tau} \equiv f_{\nu_\tau} \end{pmatrix}$

$\propto \langle a_j^\dagger(p,t)\, a_i(p,t)\rangle$

off-diagonals to take into account coherency in the neutrino system

ϱ evolution from $x H \frac{\mathrm{d}\varrho(y,x)}{\mathrm{d}x} = -ia[\mathcal{H}_{\rm eff}, \varrho] + b\mathcal{I}$

H Hubble factor → expansion (depends on universe content)

effective Hamiltonian $\mathcal{H}_{\rm eff} = \frac{\mathbb{M}_{\rm F}}{2y} - \frac{2\sqrt{2}G_{\rm F}\, y m_e^6}{x^6}\left(\frac{\mathbb{E}_\ell + \mathbb{P}_\ell}{m_W^2} + \frac{4}{3}\,\frac{\mathbb{E}_\nu}{m_Z^2}\right)$

vacuum oscillations ← → matter effects

$\mathcal{I}$ collision integrals

take into account ν–e scattering and pair annihilation, ν–ν interactions

2D integrals over momentum, take most of the computation time

solve together with z evolution, from $x\frac{\mathrm{d}\rho(x)}{\mathrm{d}x} = \rho - 3P$

ρ, P total energy density and pressure, also take into account FTQED corrections

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`FORTran-Evolved PrimordIAl Neutrino Oscillations`
`(FortEPiaNO)`
`https://bitbucket.org/ahep_cosmo/fortepiano_public`

vacuum oscillations ←

→ matter effects

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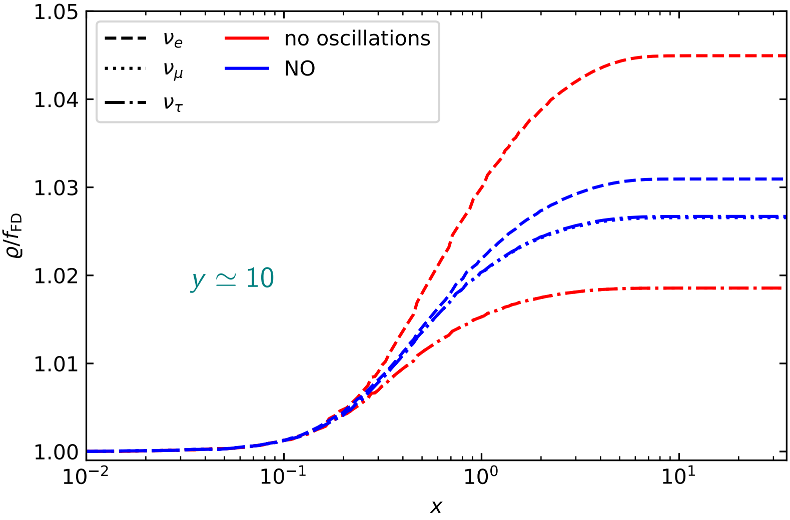
Neutrino momentum distribution and $N_{\rm eff}$ [Bennett, SG+, JCAP 2021]

Distortion of the momentum distribution ($f_{\rm FD}$: Fermi-Dirac at equilibrium)



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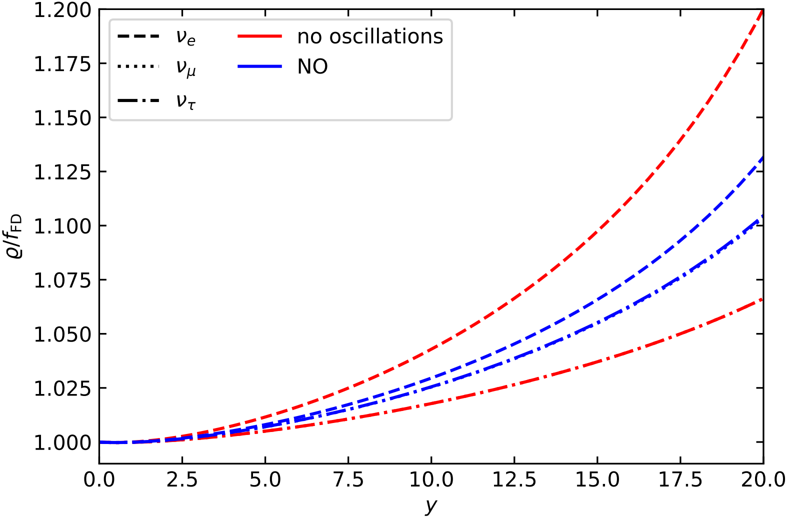
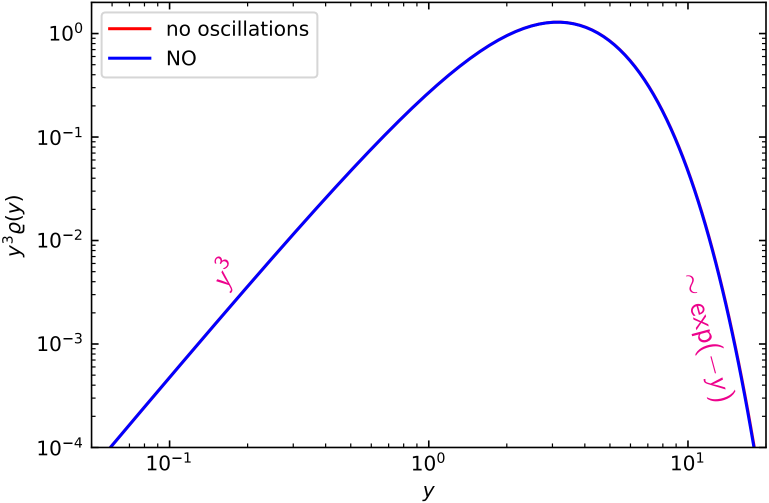


Neutrino momentum distribution and $N_{\rm eff}$ [Bennett, SG+, JCAP 2021]

$$N_{\rm eff}^{\rm final} = \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\frac{\rho_\nu}{\rho_\gamma} = \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\frac{1}{\rho_\gamma}\sum_i g_i\int\frac{d^3p}{(2\pi)^3}E(p)\,f_{\nu,i}(p)$$

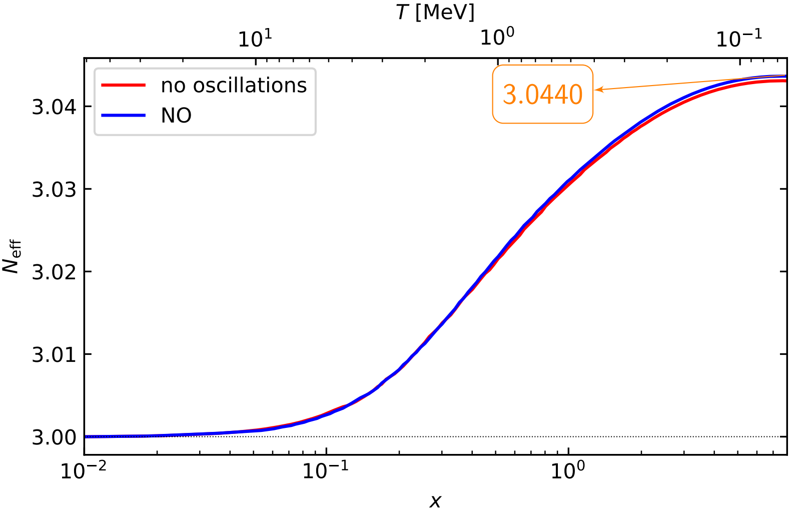
$(11/4)^{1/3} = (T_\gamma/T_\nu)^{\rm fin}$

$\hookrightarrow \propto y^3\varrho_{ii}(y)$



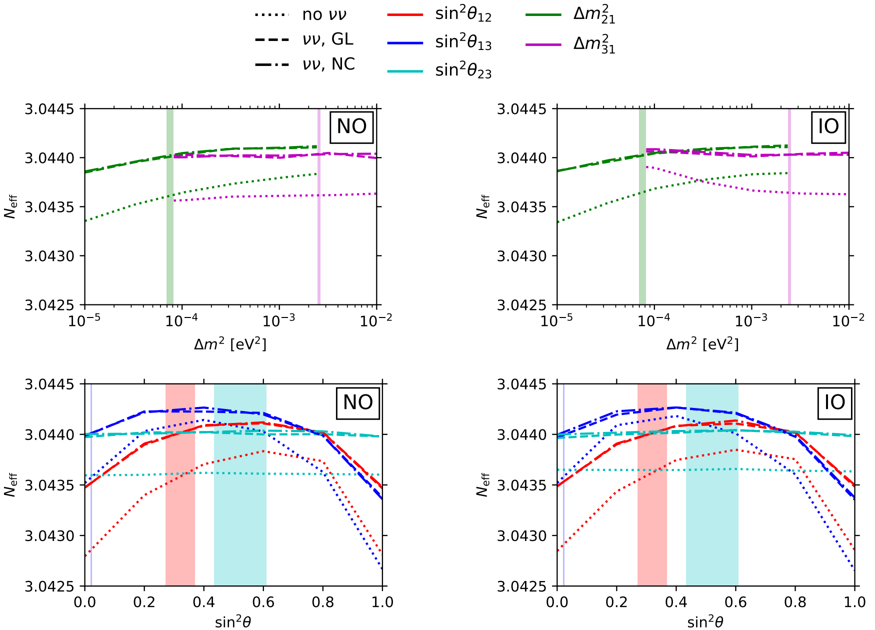
Neutrino momentum distribution and $N_{\rm eff}$ [Bennett, SG+, JCAP 2021]

$$N_{\rm eff}^{\rm any\ time} = \frac{8}{7}\left(\frac{T_\gamma}{T_\nu}\right)^4 \frac{\rho_\nu}{\rho_\gamma} = \frac{8}{7}\left(\frac{T_\gamma}{T_\nu}\right)^4 \frac{1}{\rho_\gamma}\sum_i g_i \int \frac{d^3p}{(2\pi)^3} E(p)\, f_{\nu,i}(p)$$



Effect of neutrino oscillations

[Bennett, SG+, JCAP 2021]



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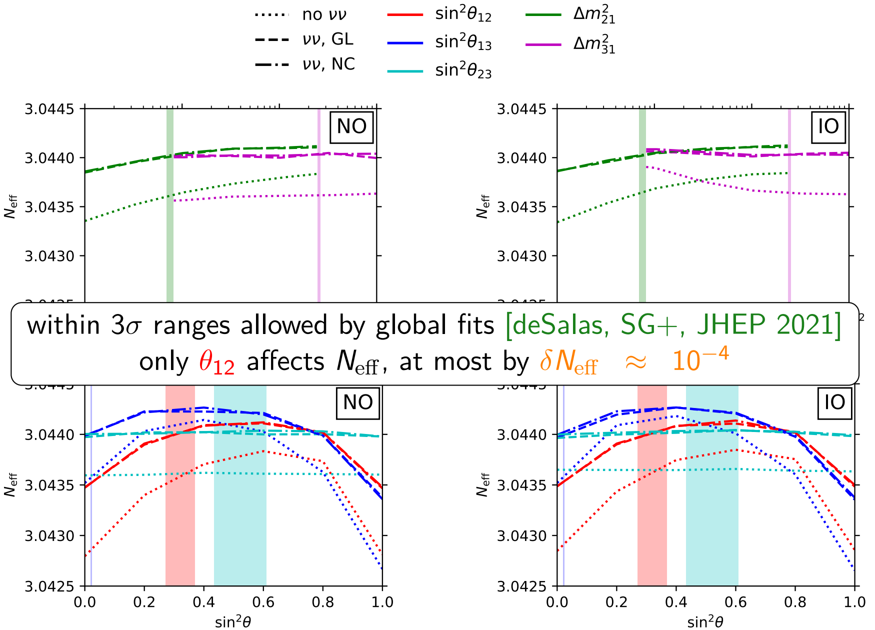
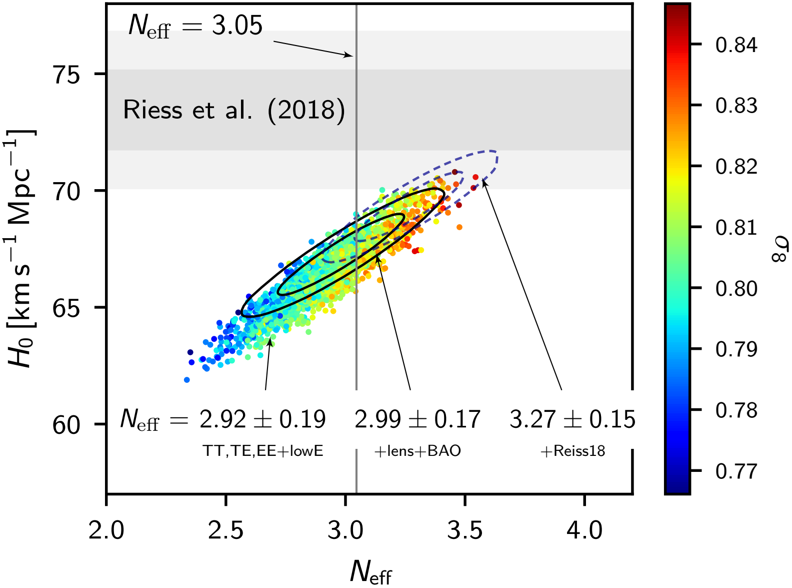


within 3σ ranges allowed by global fits [deSalas, SG+, JHEP 2021]
only θ_{12} affects $N_{\rm eff}$, at most by $\delta N_{\rm eff} \approx 10^{-4}$

$N_{\rm eff}$ and CMB

[Planck Collaboration, 2018]



$N_{\rm eff}$ and BBN

BBN: production of light nuclei at $t \sim 1$s to $t \sim \mathcal{O}(10^2)$s

temperature $T_{fr} \simeq 1$ MeV
from nucleon freeze-out:

$$\Gamma_{n\leftrightarrow p} \sim G_F^2 T^5 \;=\; H \sim \sqrt{g_\star G_N}\, T^2$$

$$\downarrow$$

$$\boxed{T_{fr} \simeq (g_\star G_N/G_F^4)^{1/6}}$$

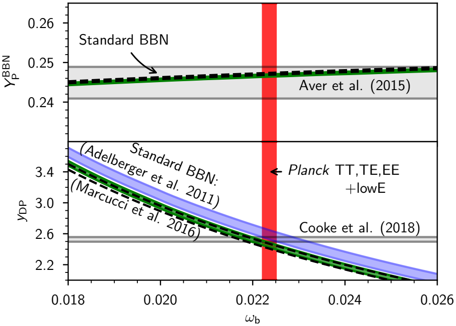
enters
$n/p = \exp(-Q/T_{fr})$
which controls element abundances

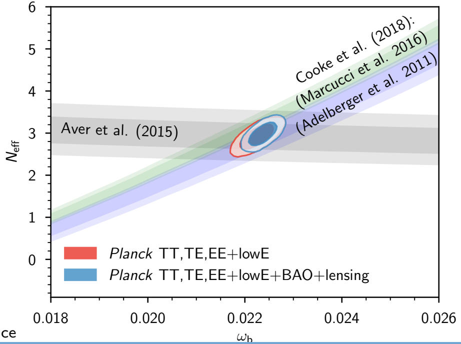
$\boxed{g_\star \text{ depends on } N_{\rm eff}}$

$$\downarrow$$

abundances depend on $N_{\rm eff}$

[Planck Collaboration, 2018]





G_F Fermi constant
G_N Newton constant

n, p: neutron, proton density number
$Q = 1.293$ MeV neutron–proton mass difference

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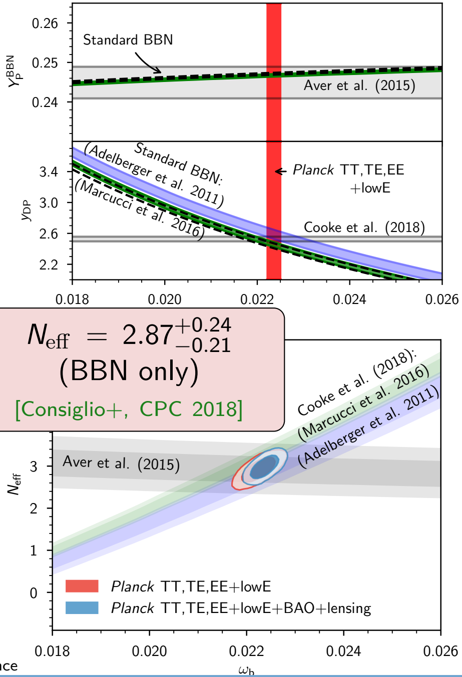

[Planck Collaboration, 2018]

$$N_{\rm eff} = 2.87^{+0.24}_{-0.21}$$

(BBN only)

[Consiglio+, CPC 2018]

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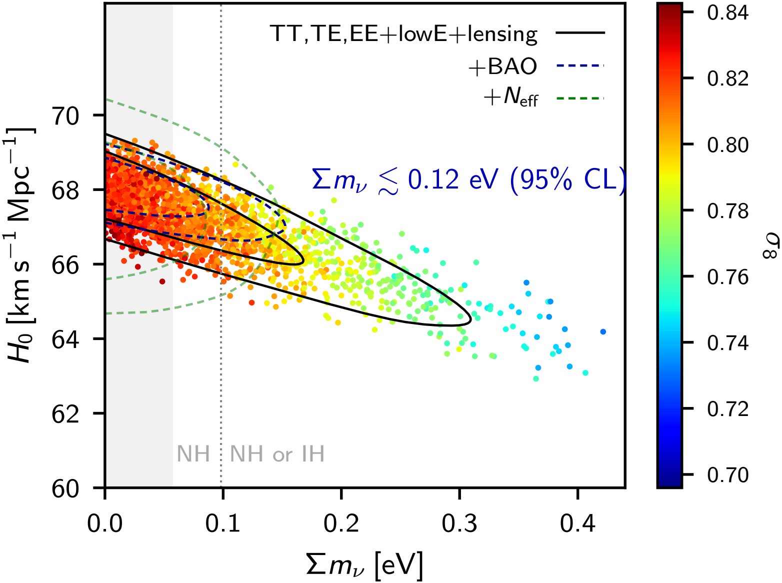
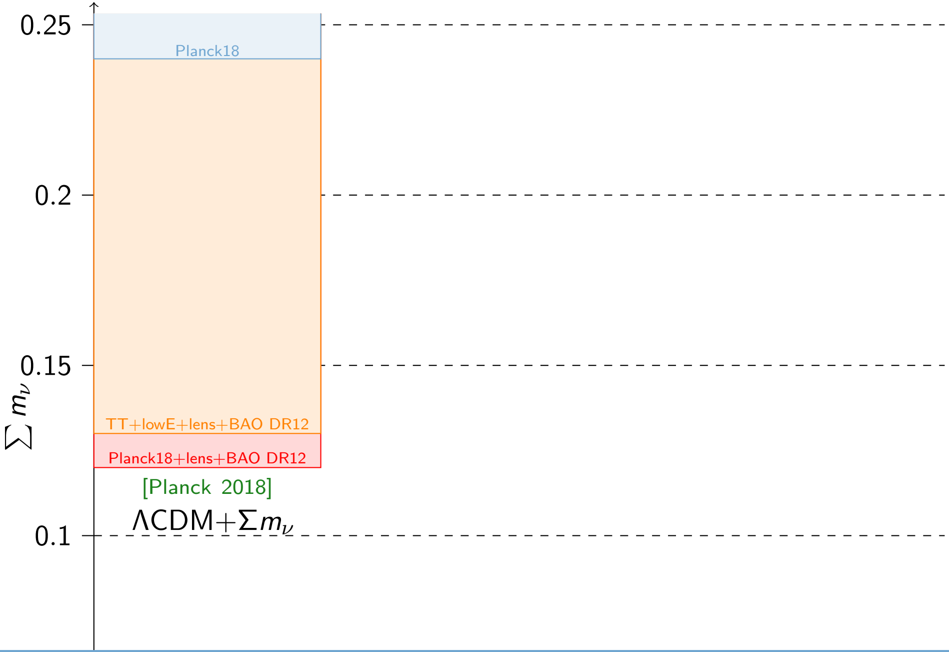
Σm_ν and CMB

[Planck Collaboration, 2018]



Cosmological neutrino mass bounds (95% CL)



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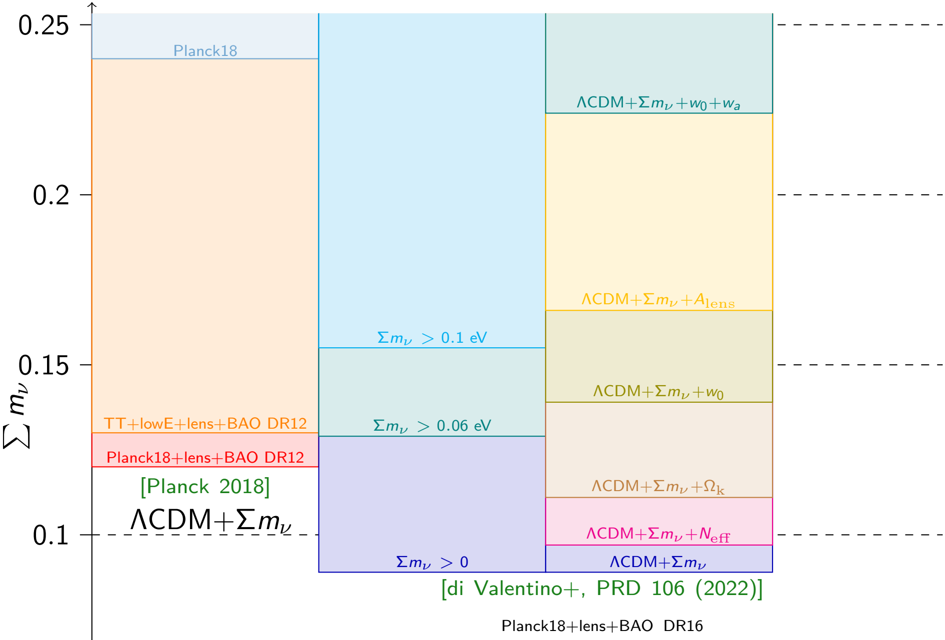
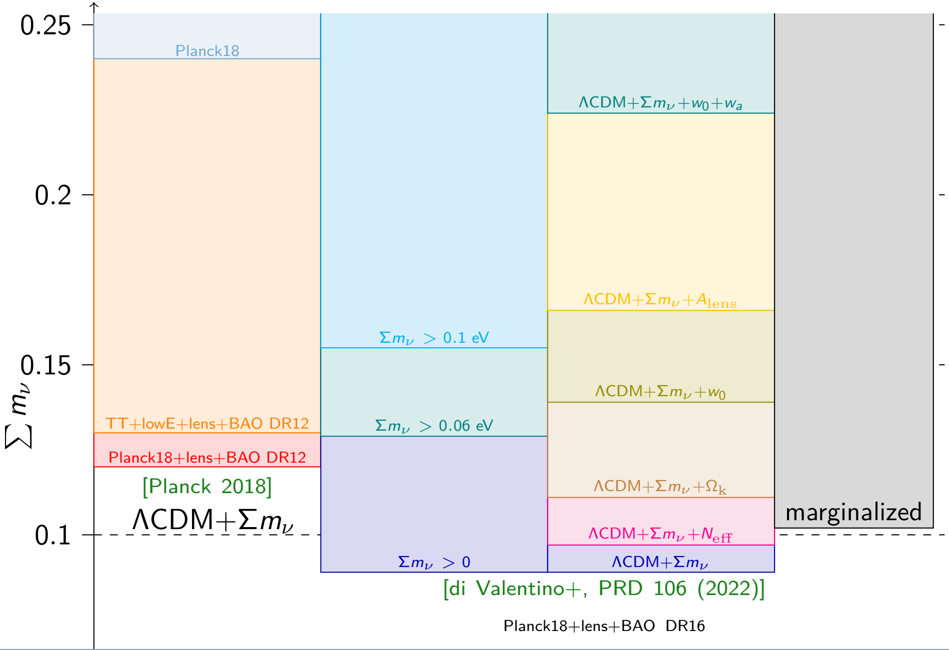


Cosmological neutrino mass bounds (95% CL)



Can a cosmological limit on Σm_ν disfavor IO?

[PDU 40 (2023)]

standard factor

Cosmology measures $\omega_\nu = \Omega_\nu h^2 = \Sigma m_\nu / (94.12 \text{ eV})$

NO: $\Sigma m_\nu \gtrsim 0.06$ eV

Current: $\Sigma m_\nu \lesssim 0.1$ eV (95%)

IO: $\Sigma m_\nu \gtrsim 0.1$ eV

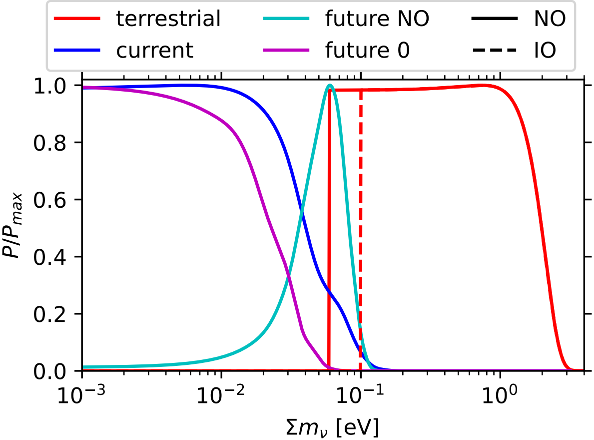
Future sensitivity: $\sigma(\Sigma m_\nu) \simeq 0.02$ eV

Still preferring $\Sigma m_\nu = 0$?

tension ever with NO!

Will measure e.g. $\Sigma m_\nu = 0.06$ eV?

confirm NO, disfavor IO



Can a cosmological limit on Σm_ν disfavor IO?

[PDU 40 (2023)]

standard factor

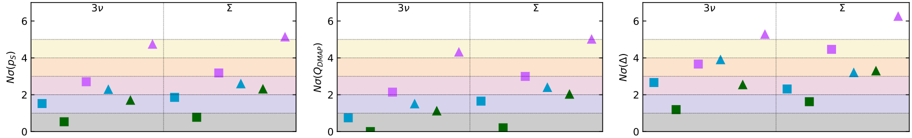
Cosmology measures $\omega_\nu = \Omega_\nu h^2 = \Sigma m_\nu/(94.12\text{ eV})$

Is there a tension between cosmology and oscillations?

or will there be a tension?

several possible tests can be considered, similar results

$\Sigma m_\nu \lesssim 0.1$ eV (95%) ● current ■ NO
$\Sigma m_\nu = 0.06 \pm 0.02$ eV (1σ) ● future NO ▲ IO
$\Sigma m_\nu = 0.00 \pm 0.02$ eV (1σ) ● future 0



currently only mild tension between cosmology and oscillations

future NO can be at $\sim 2\sigma$ tension with IO

future 0 can be at $\sim 2-3\sigma$ tension with NO, $\gtrsim 4\sigma$ with IO

Can a cosmological limit on Σm_ν disfavor IO?

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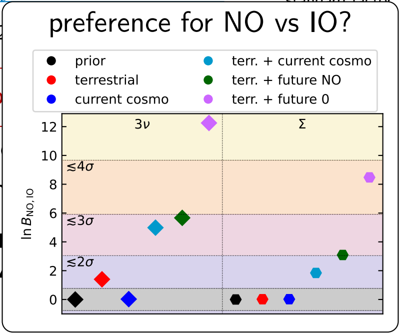
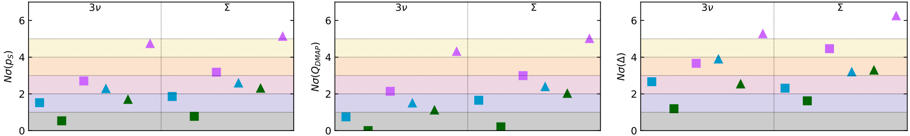


preference for NO vs IO?



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future 0 can be at $\sim 2-3\sigma$ tension with NO, $\gtrsim 4\sigma$ with IO

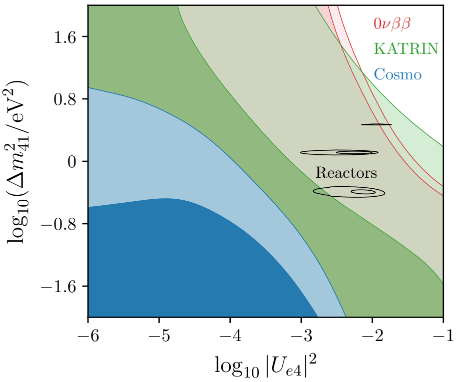
B

Sterile neutrinos

let's pretend they exist

Based on:

- JPG 43 (2016) 033001
- JHEP 06 (2017) 135
- PLB 782 (2018) 13-21
- JCAP 07 (2019) 014
- PRD 104 (2021) 123524
- JCAP 03 (2023) 046



Short Baseline (SBL) anomalies

[SG+, JPG 43 (2016) 033001]

Do three-neutrino oscillations explain all experimental results?

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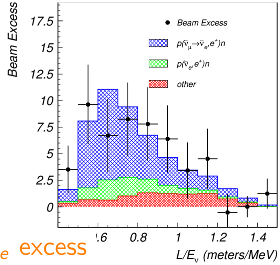
LSND



[Aguilar+, 2001]

3.8σ

$\bar{\nu}_\mu \to \bar{\nu}_e$ excess

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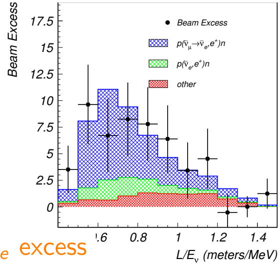

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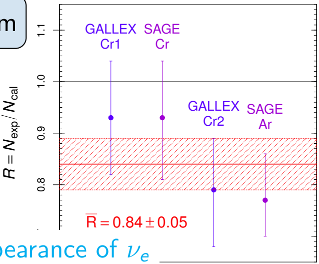
Gallium



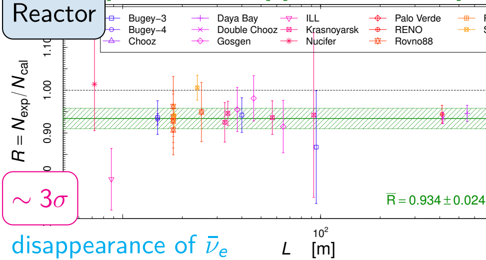
[Giunti, Laveder, 2011]

2.7σ

disappearance of ν_e

[Mention+, 2011] [Azabajian+, 2012]

Reactor



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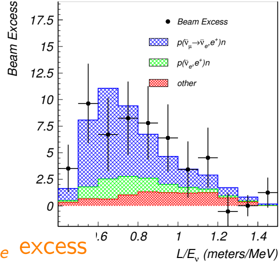

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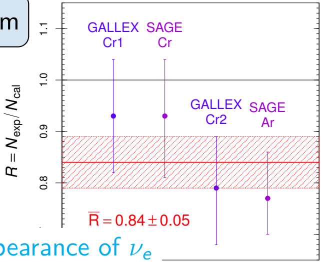

[Giunti, Laveder, 2011]

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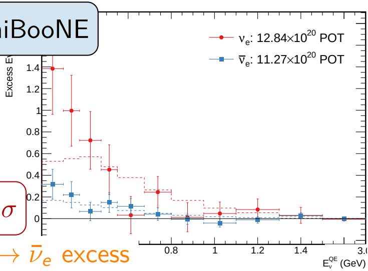
MiniBooNE



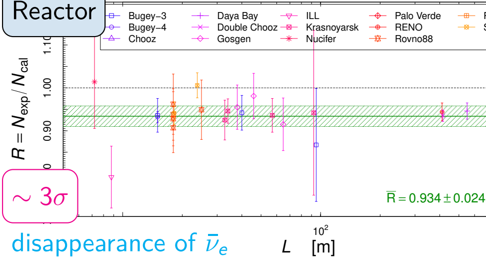
[Aguilar+, 2008-2018]

$\sim 5\sigma$

$\bar{\nu}_\mu \to \bar{\nu}_e$ excess

Reactor

[Mention+, 2011] [Azabajian+, 2012]



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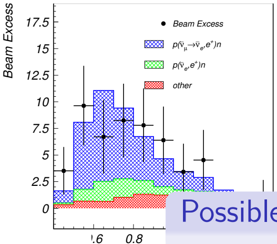

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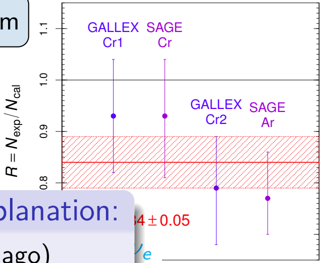


[Giunti, Laveder, 2011]

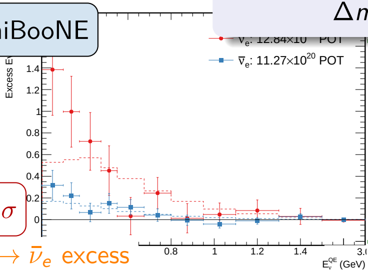
Possible common explanation:

(until a few years ago)

Additional squared mass difference

$\Delta m^2_{\rm SBL} \simeq 1\ {\rm eV}^2$

MiniBooNE



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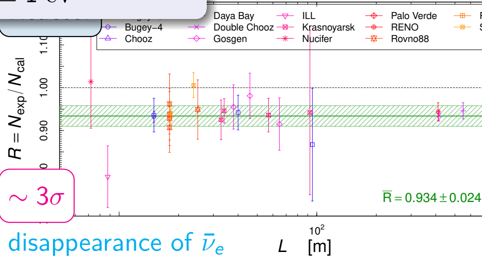

$\sim 3\sigma$

disappearance of $\bar{\nu}_e$

Short BaseLine (SBL)

$$P_{\nu_\alpha\to\nu_\beta}(L) = |\langle\nu_\alpha|\nu(L)\rangle|^2 = \sum_{k,j} U_{\beta k} U^*_{\alpha k} U^*_{\beta j} U_{\alpha j} \exp\left(-i\frac{\Delta m^2_{kj} L}{2E}\right)$$

If $m_4 \gg m_\ell$, faster oscillations

ν_4 oscillations are averaged in most neutrino oscillation experiments

Effect of 4th neutrino only visible as global normalization

Short BaseLine (SBL) oscillations: $\frac{\Delta m^2_{41} L}{E} \simeq 1$

At SBL, oscillations due to Δm^2_{21} and $|\Delta m^2_{31}|$ do not develop

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source of $\overset{(-)}{\nu}_\alpha$

detect different flavor $\overset{(-)}{\nu}_\beta$

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DISappearance

source of $\overset{(-)}{\nu}_\alpha$

$\overset{(-)}{\nu}_s$ goes undetected

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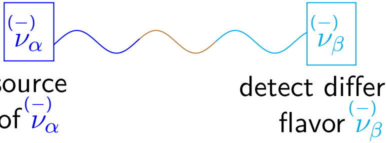
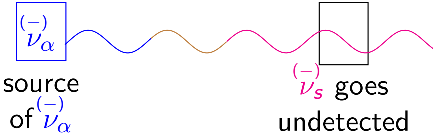
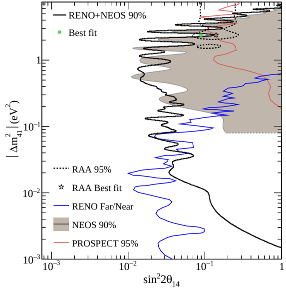




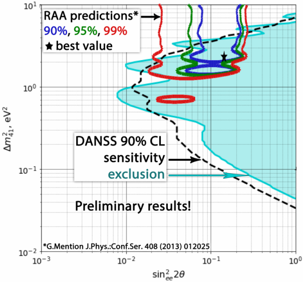
CP violation cannot be observed in SBL experiments!

ν_s at reactors in 2020

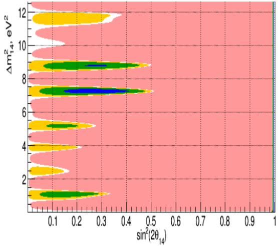
[RENO+NEOS, 2020]



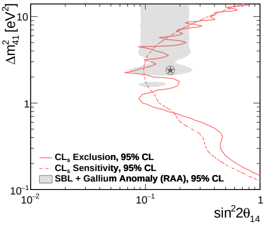
[DANSS, 2020]



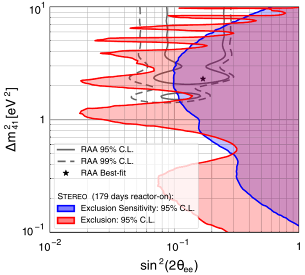
[Neutrino-4, PZETF 2020]



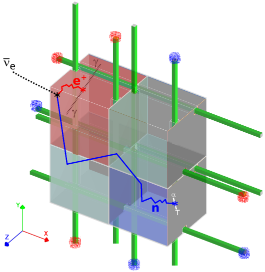
[PROSPECT, PRD 2020]



[STEREO, PRD 2020]

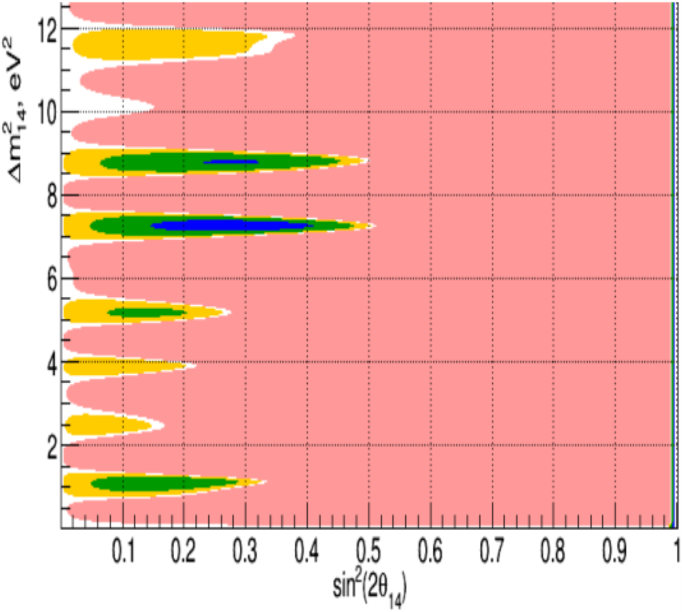


[SoLiD, JINST 2021]



Neutrino-4

[Neutrino-4 2020]

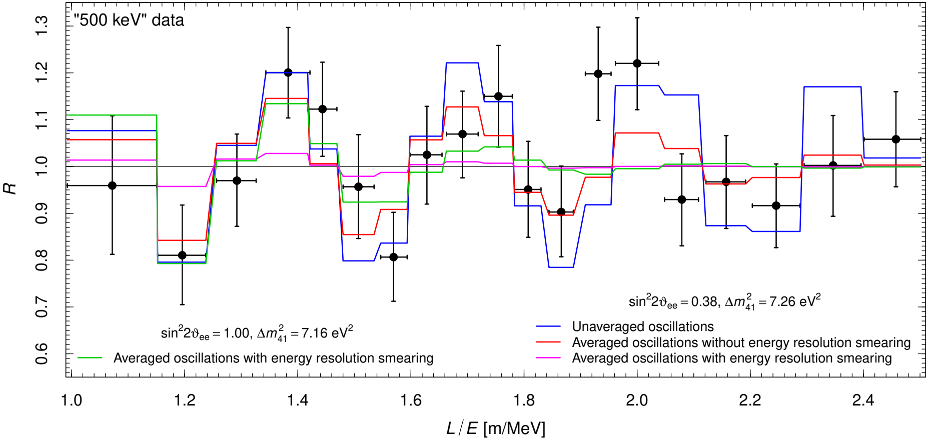


claimed $> 3\sigma$ preference for 3+1 over 3ν case

best fit incompatible with other reactor experiments

Neutrino-4

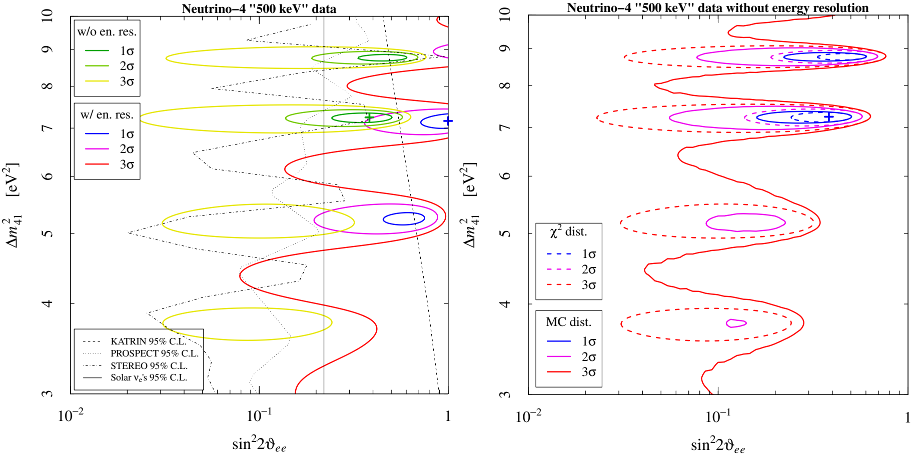
[Giunti+, PLB 2021]



energy resolution smearing not properly taken into account?

Neutrino-4

[Giunti+, PLB 2021]



proper energy resolution treatment moves best-fit → $\sin^2 2\vartheta \simeq 1$

need to take into account violation of Wilk's theorem

↓

relaxed constraints

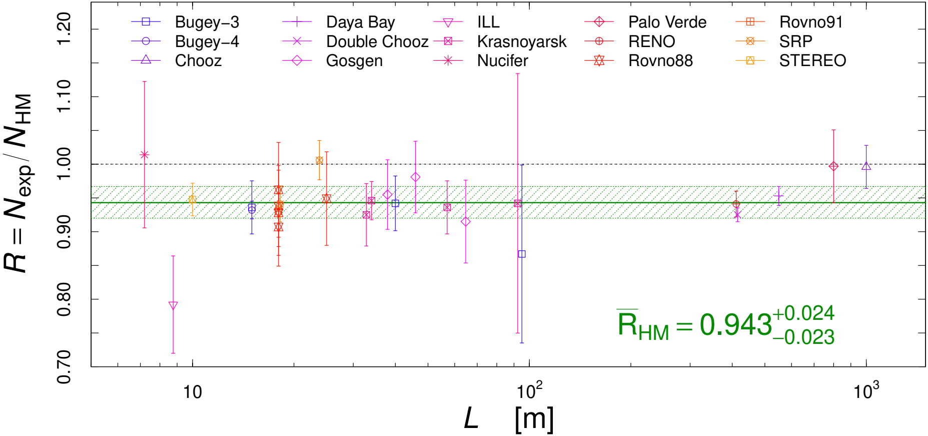
Reactor antineutrino spectrum and RAA

[Giunti+, PLB 2022]

When the RAA was discovered:

conversion method (ILL data) and *ab initio* calculations in agreement

[Huber, 2011], [Mueller+, 2011] spectra



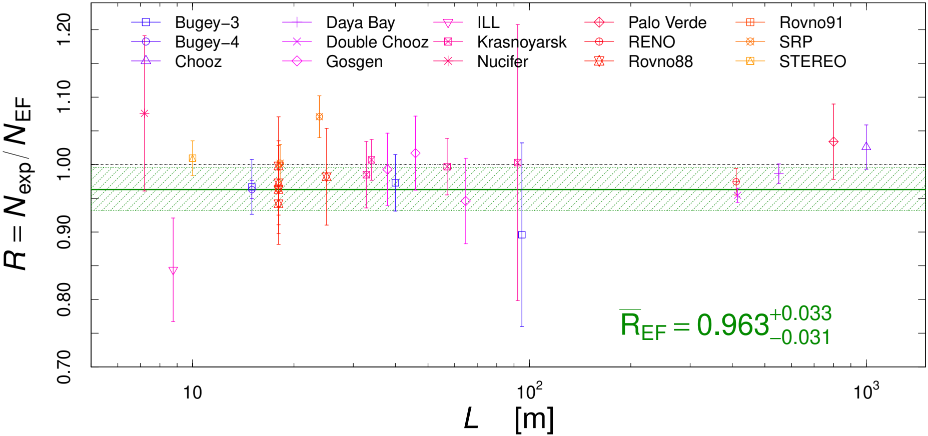
$\sim 2.4\sigma$ deficit $\Longrightarrow$ anomaly!

Reactor antineutrino spectrum and RAA

[Giunti+, PLB 2022]

Revised *ab initio* calculation:

[Estienne, Fallot+, PRL 123 (2019)]



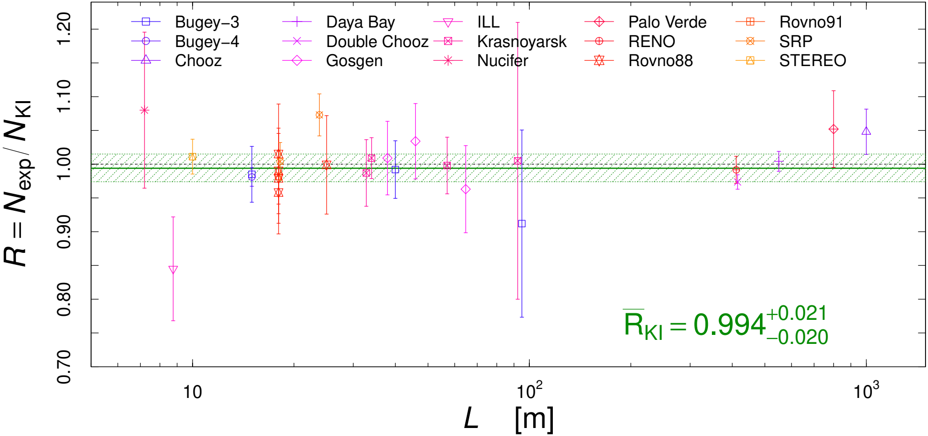
$\sim 1.2\sigma$ deficit $\Longrightarrow$ no anomaly!

Reactor antineutrino spectrum and RAA

[Giunti+, PLB 2022]

Conversion method on new measurements of electron spectrum at Kurchatov Institute (KI) (updates ILL measurements from the 80's):

[Kopeikin+, PRD 2021]



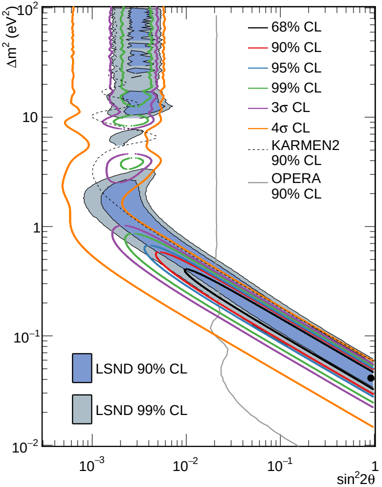
approximate agreement with EF fluxes, no anomaly!

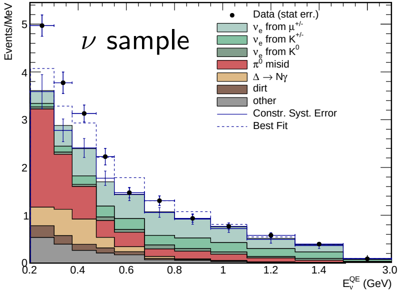
MiniBooNE

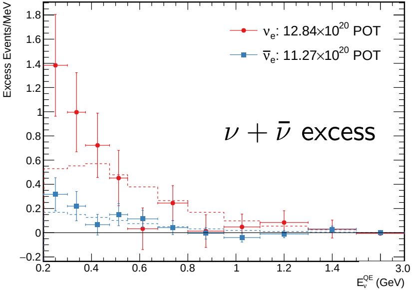
purpose: check LSND signal

$L \simeq 541$ m, 200 MeV $\leq E \lesssim 3$ GeV

no money, no near detector







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[PRL 121 (2018) 221801]

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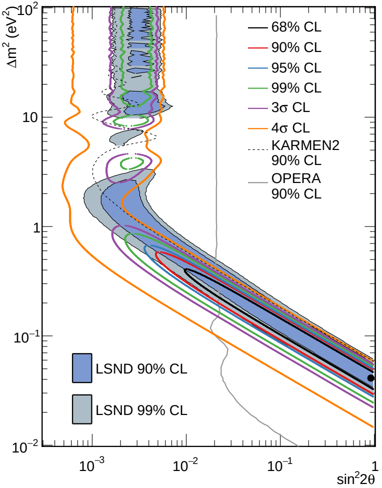
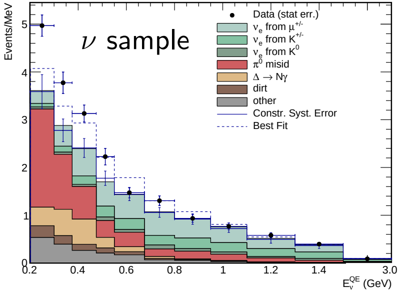
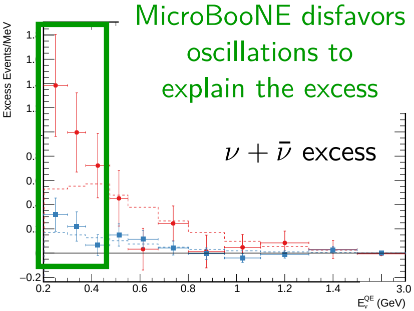


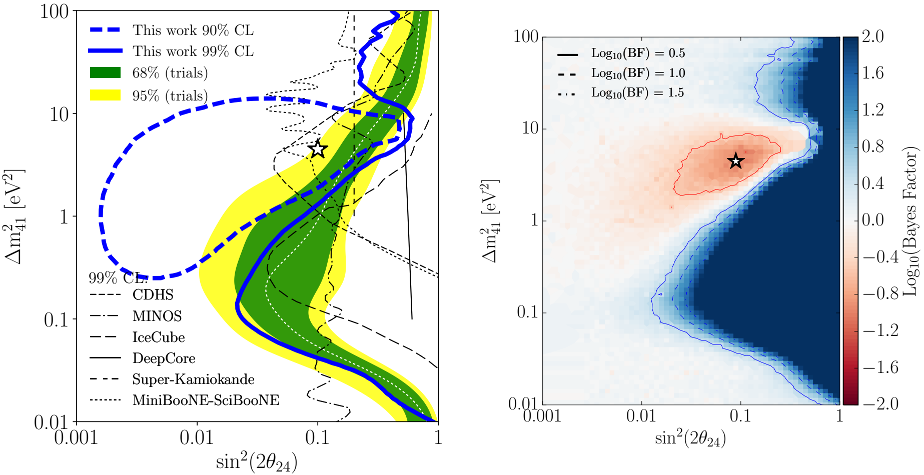




MicroBooNE disfavors oscillations to explain the excess

IceCube 8 yr update

[IceCube, PRL 2020]



first indication in favor of sterile from ν_μ DIS!

although rather weak: $\log_{10} BF \simeq 1$ (weak preference)

or compatible with no oscillations at p-value of 8%

Sterile neutrino in the early universe

[SG+, JCAP 07 (2019) 014]

Four neutrinos ⟶ new oscillations in the early Universe

sterile ⟹ no weak/em interactions in the thermal plasma

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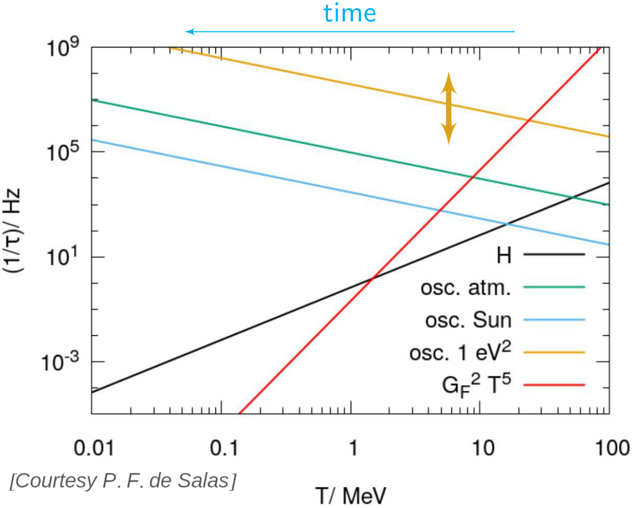
need to produce it through oscillations, but matter effects may block them



[Courtesy P. F. de Salas]

beginning of oscillations depends on Δm^2_{41}

later oscillations

⇓

less time before ν decoupling!

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when are they enough to allow full equilibrium of active-sterile states?

$$0 \longleftarrow \Delta N_{\rm eff} = N_{\rm eff}^{4\nu} - N_{\rm eff}^{3\nu} \longrightarrow \simeq 1$$

no sterile production

active&sterile in equilibrium

$$\frac{\Delta m^2_{as}}{\mathrm{eV}^2}\sin^4(2\vartheta_{as}) \simeq 10^{-5}\ln^2(1-\Delta N_{\rm eff}) \qquad \text{(1+1 approx.)}$$

[Dolgov&Villante, 2004]

e.g.: $\Delta m^2_{as} = 1\ \mathrm{eV}^2$, $\sin^2(2\vartheta_{as}) \simeq 10^{-3} \Longrightarrow \Delta N_{\rm eff} \simeq 1$

$N_{\rm eff}^{3\nu} = 3.044$ [JCAP 2021]

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Full calculation: use numerical code!

`FORTran-Evolved PrimordIAl Neutrino Oscillations`

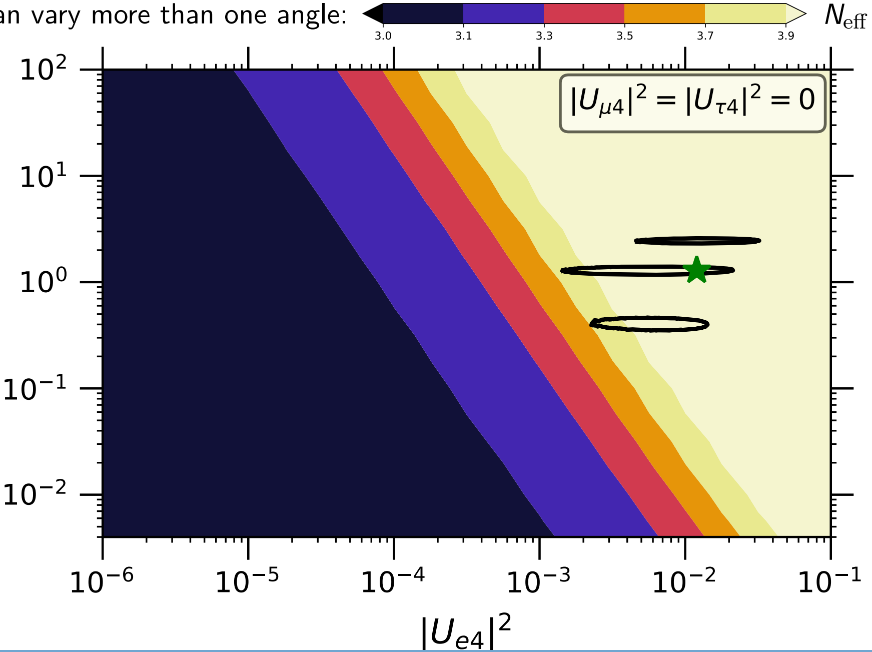
`(FortEPiaNO)`

`https://bitbucket.org/ahep_cosmo/fortepiano_public`

$N_{\rm eff}$ and the new mixing parameters

[SG+, JCAP 07 (2019) 014]

We can vary more than one angle:



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[SG+, JCAP 07 (2019) 014]

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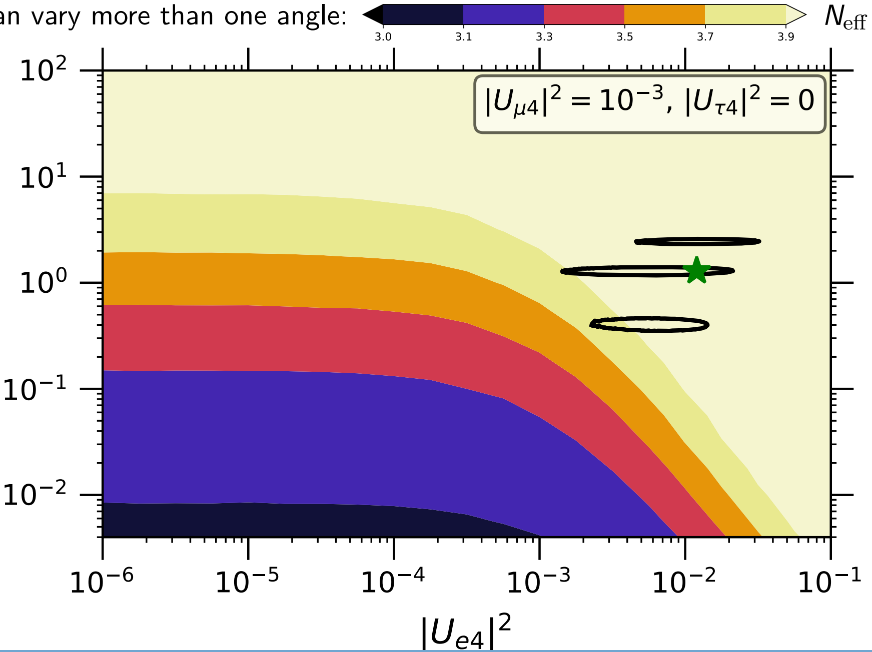

$N_{\rm eff}$ and the new mixing parameters

[SG+, JCAP 07 (2019) 014]

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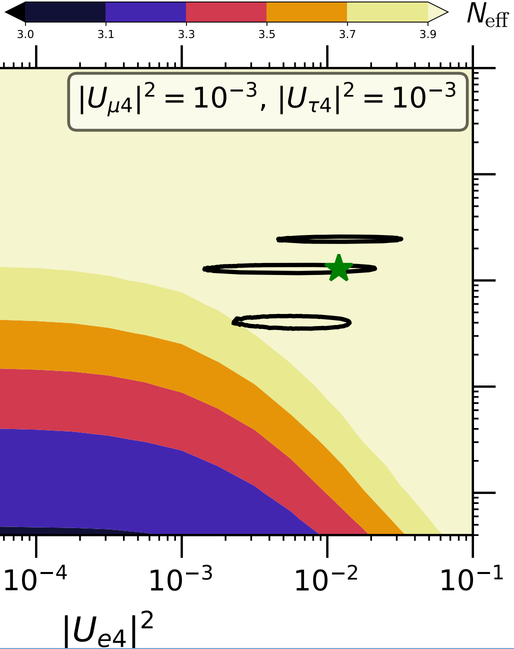
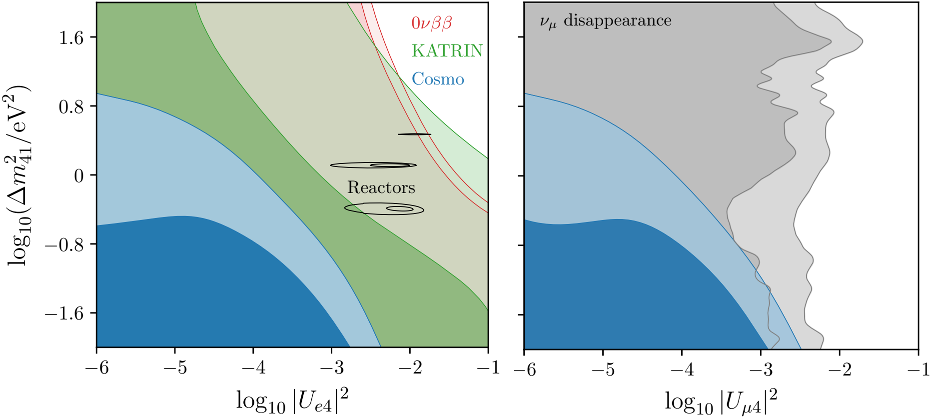


Comparing constraints

[PRD 104 (2021) 123524]

Cosmological constraints are stronger than most other probes

But much more model dependent (as all the cosmological constraints)!



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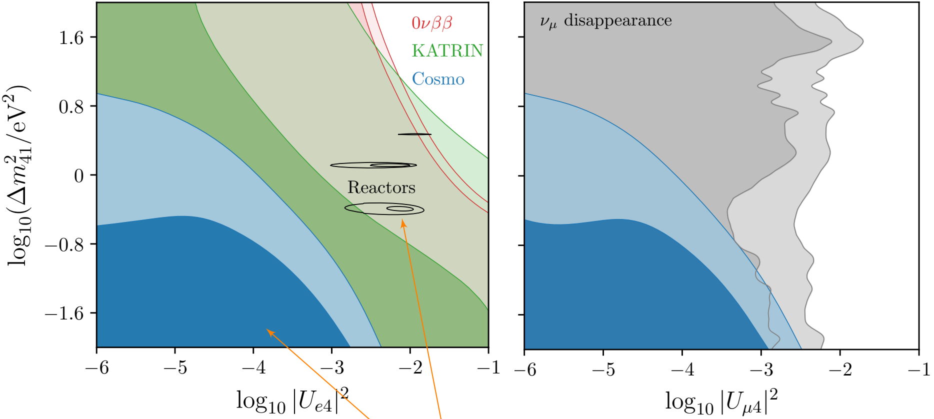

Warning: tension between reactor experiments and CMB bounds!

Comparing constraints

[PRD 104 (2021) 123524]

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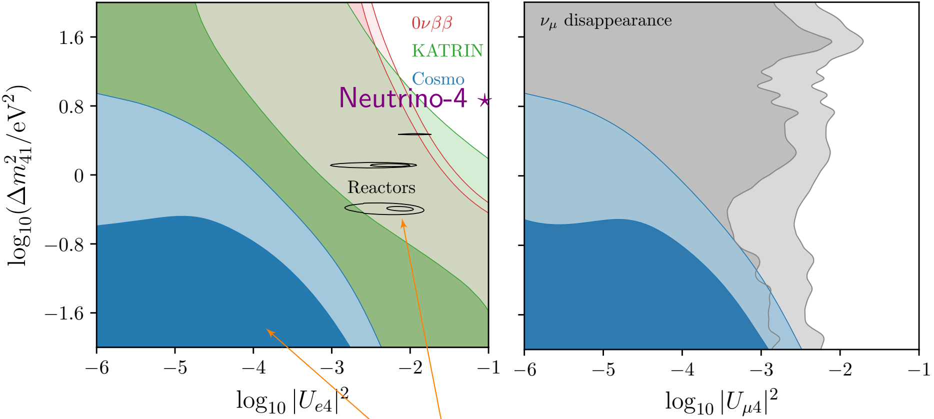

Warning: tension between reactor experiments and CMB bounds!

Non-unitarity of the 3 × 3 mixing matrix

[JCAP 03 (2023)]

Consider we have N_ν neutrino states

Unitary $N_\nu \times N_\nu$ mixing matrix: $V = \begin{pmatrix} \boxed{\begin{matrix} V_{e1} & V_{e2} & V_{e3} \\ V_{\mu1} & V_{\mu2} & V_{\mu3} \\ V_{\tau1} & V_{\tau2} & V_{\tau3} \end{matrix}} & \cdots \\ \vdots & \ddots \end{pmatrix}$

the 3×3 sector (N)
describing mixing among lightest neutrinos
is non-unitary

$$N = \begin{pmatrix} \alpha_{11} & 0 & 0 \\ \alpha_{21} & \alpha_{22} & 0 \\ \alpha_{31} & \alpha_{32} & \alpha_{33} \end{pmatrix} U$$

α_{ii} real, α_{ij} $(i \neq j)$ complex $\Rightarrow$ CP violation

$U = R^{23} R^{13} R^{12}$ is the standard unitary mixing matrix

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[JCAP 03 (2023)]

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the 3×3 sector (N)
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Neutrino interactions depend only on kinematically accessible states

Oscillations depend on all states

Oscillations with states $n > 3$ much heavier than $n \leq 3$

are averaged out at experiments

Non-unitarity and neutrino decoupling

[JCAP 03 (2023)]

Neutrino density matrix evolution in mass basis:

$$\left.\frac{\mathrm{d}\varrho(y)}{\mathrm{d}x}\right|_{\mathrm{M}} = \sqrt{\frac{3m_{\mathrm{Pl}}^2}{8\pi\rho}} \left\{ -i\frac{x^2}{m_e^3}\left[\frac{\mathbb{M}_{\mathrm{M}}}{2y} - \frac{2\sqrt{2}G_F y m_e^6}{x^6}\mathcal{E}_{\mathrm{M}}, \varrho\right] + \frac{m_e^3}{x^4}\mathcal{I}(\varrho)\right\}$$

Unitary case

interactions:

$$\begin{aligned}(Y_L)_{ab} &\equiv \tilde{g}_L\mathbb{I} + (U^\dagger)_{ea}U_{eb}\\ (Y_R)_{ab} &\equiv g_R\mathbb{I}\end{aligned}$$

matter effects:

$$\mathcal{E}_{\mathrm{M}} = \frac{\rho_e+P_e}{m_W^2}U^\dagger \mathrm{diag}(1,0,0)U$$

Fermi constant:

$$G_F^\mu = G_F$$

Non-unitary case

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$$\begin{aligned}(Y_L)_{ab} &\equiv \tilde{g}_L(V^\dagger V)_{ab} + (V^\dagger)_{ea}V_{eb}\\ (Y_R)_{ab} &\equiv g_R(V^\dagger V)_{ab}\end{aligned}$$

matter effects:

$$\mathcal{E}_{\mathrm{NU}} \equiv \frac{\rho_e+P_e}{m_W^2}(Y_L - Y_R)$$

Fermi constant:

$$G_F^\mu = G_F\sqrt{\alpha_{11}^2\left(\alpha_{22}^2 + |\alpha_{21}|^2\right)}$$

$$G_F^\mu = 1.1663787(6)\times 10^{-5}\ \mathrm{GeV}^{-2}$$ [CODATA]

$$\mathcal{I}(\varrho) \propto G_F^2$$

Non-unitarity parameters and $N_{\rm eff}$

[JCAP 03 (2023)]

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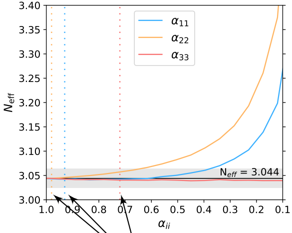
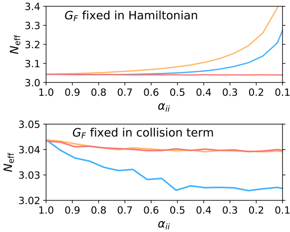


terrestrial bounds



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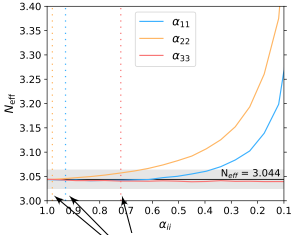
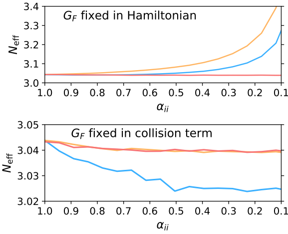
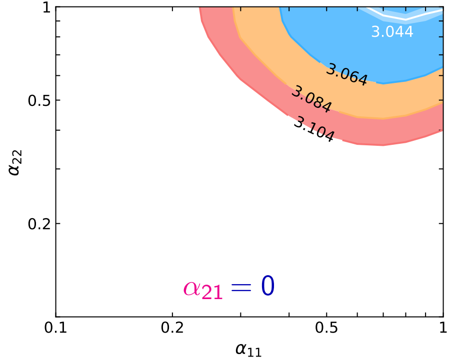
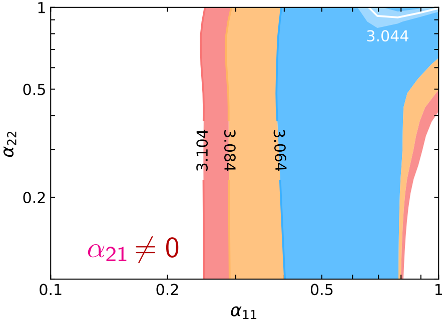








(α_{21} marginalized over)

Confidence regions from future CMB measurements with $\delta N_{\rm eff} = 0.02$

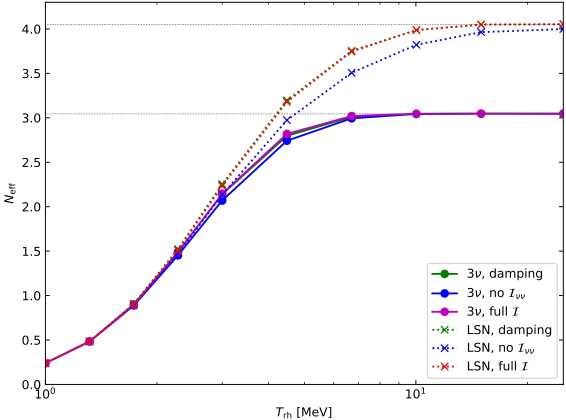
C

Non-standard cosmology

beyond non-standard neutrino physics

Based on:

- arxiv:2308.15531 (1st secondment)
- in preparation (2nd secondment)



Additional particles in the early universe?

[arxiv:2308.15531]

Sterile neutrinos are coupled via oscillations to the thermal plasma
(photons, electrons, neutrinos, (muons), ...)

What if we add a decoupled particle?

let us assume a non-standard evolution of the energy density: $\bar{\rho}_{\rm US} \propto a^{n+4}$
$n = 0 \rightarrow$ radiation; $n = -1 \rightarrow$ matter; $n = -2 \rightarrow$ curvature, ...

effect on early universe phenomena is purely gravitational

total energy density: $\rho = \rho_\gamma + \rho_e + \rho_\nu + \delta\rho_{\rm FTQED} + \rho_{\rm US}$

Hubble factor: $H^2 = 8\pi\rho/(3M_{\rm Pl}^2)$

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neutrino decoupling: $$\frac{\mathrm{d}\varrho(y)}{\mathrm{d}x} = \frac{1}{xH}\left\{-i\frac{x^3}{m_e^3}\left[\mathcal{H}_{\rm eff}, \varrho\right] + \frac{m_e^3}{x^3}\mathcal{I}(\varrho)\right\}$$

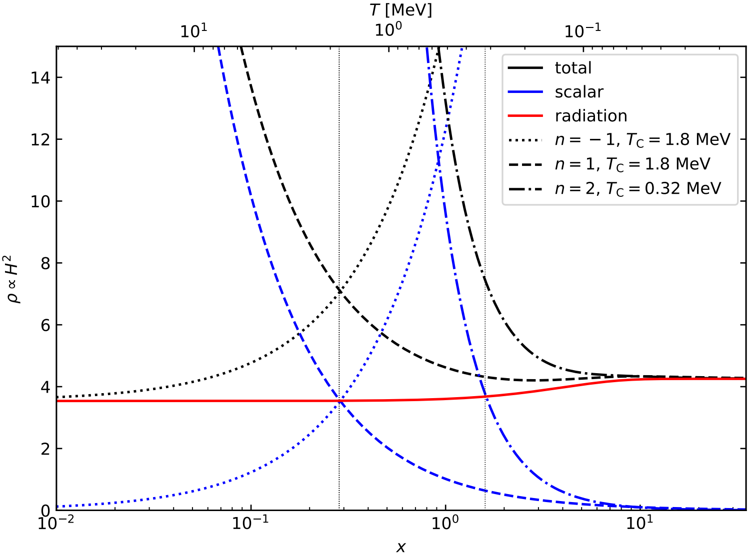
BBN abundances: $$\frac{dX_i}{dx} = \frac{\Gamma_i}{xH}$$

$X_i = n_i/N_B$ abundance relative to total baryons, Γ_i effective reaction rate for nuclide i

Results from $N_{\rm eff}$

consider $\rho_{\rm US} = \rho_{\rm rad}$ at $x_{\rm C} = m_e/T_{\rm C}$ for the new particle

Evolution of the energy density:



Results from $N_{\rm eff}$

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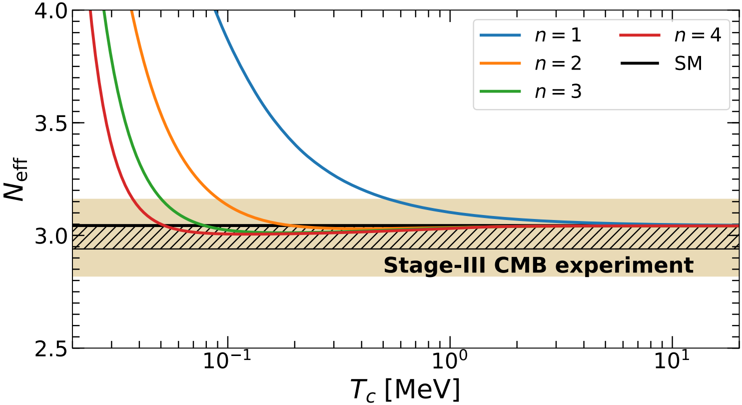
From neutrino decoupling we obtain:

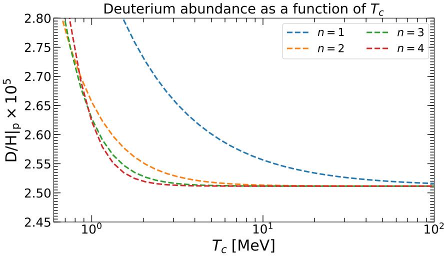


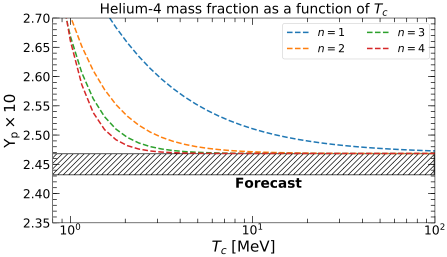
Results from BBN

[arxiv:2308.15531]

consider $\rho_{\rm US} = \rho_{\rm rad}$ at $x_{\rm C} = m_e/T_{\rm C}$ for the new particle

Compare to current measurements (Deuterium, Helium):


Deuterium abundance as a function of T_c


Helium-4 mass fraction as a function of T_c

error bands (gray) are current constraints on the abundances

even current precision can strongly constrain $T_{\rm C}$

calculations performed during and after the secondment,
with D. Aristizabal and A. Villanueva (UTFSM, Chile)

Scenarios with low reheating temperature

[in preparation]

Reheating: phase ending inflation

during inflation, the inflaton (non-rel. scalar) dominates the energy density

during reheating: inflaton decays into standard model particles

$\Longrightarrow$ photons, electrons, ... are populated directly

radiation domination begins after reheating

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neutrinos are populated by weak interactions with electrons!

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Low reheating temperature: when reheating occurs at $T_{\rm rh} \lesssim 20$ MeV

notice: if $T_{\rm rh} \lesssim 3$ MeV, BBN is broken!

3 neutrino oscillations start to be affected when $T_{\rm rh} \lesssim 8$ MeV

what about sterile neutrinos?

$N_{\rm eff}$ with low reheating

[in preparation]

Need to edit equations for inflaton energy density and its contribution:

$$\frac{\mathrm{d}\varrho(y)}{\mathrm{d}x} = \text{unchanged}$$

$$\frac{\mathrm{d}\rho_\phi}{\mathrm{d}x} = -\frac{x\rho_\phi\Gamma_\phi}{m_e^2}\sqrt{\frac{3m_{\rm Pl}^2}{8\pi\rho_{\rm tot}}}$$

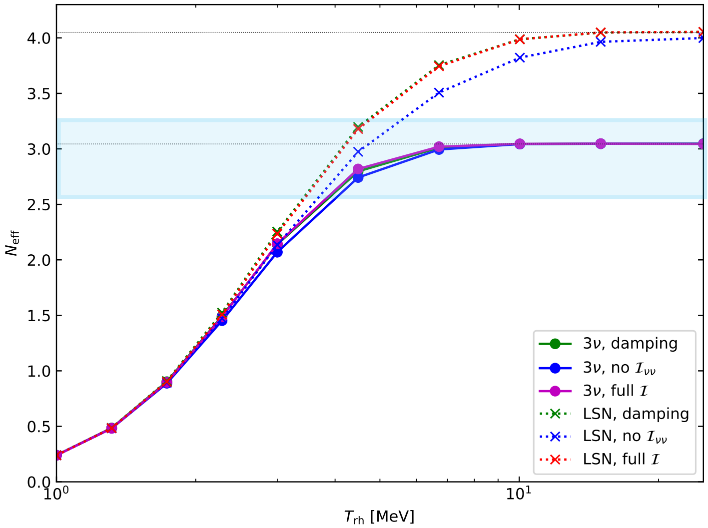
$$\frac{\mathrm{d}z}{\mathrm{d}x} = \frac{\sum_{\ell=e,\mu}\left[\frac{r_\ell^2}{r}J_2(r_\ell)\right] + G_1(r) - \frac{1}{2z^3}\sum_{\alpha=e}^{s}\frac{\mathrm{d}\rho_{\nu_\alpha}}{\mathrm{d}x} - \frac{x}{2z^3}\frac{\mathrm{d}\rho_\phi}{\mathrm{d}x}}{\sum_{\ell=e,\mu}\left[r_\ell^2 J_2(r_\ell) + J_4(r_\ell)\right] + G_2(r) + \frac{2\pi^2}{15}}$$

$$\rho_{\rm tot} = \sum_{i=\gamma,\nu_j,e,\mu}\rho_i + \delta\rho(x,z) + x\rho_\phi$$

$$\Gamma_\phi \simeq \left(\frac{T_{\rm rh}}{0.7\mathrm{MeV}}\right)^2 \sec^{-1}$$

$N_{\rm eff}$ with low reheating

$N_{\rm eff}$ as a function of $T_{\rm rh}$ (3 or 3+1 neutrinos):

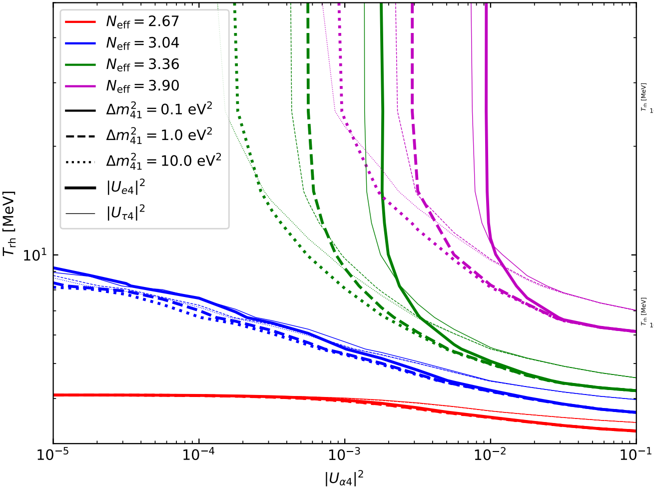


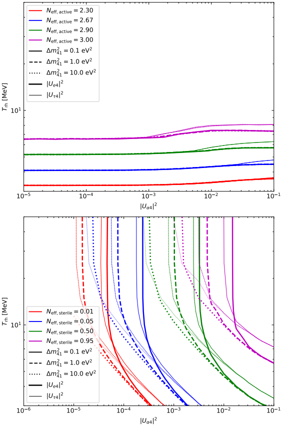
Planck constraint: $N_{\rm eff} = 2.92^{+0.36}_{-0.37}$ (95%, TT,TE,EE+lowE)

$N_{\rm eff}$ with low reheating

[in preparation]

sterile case with varying mixing angle/mass splitting:





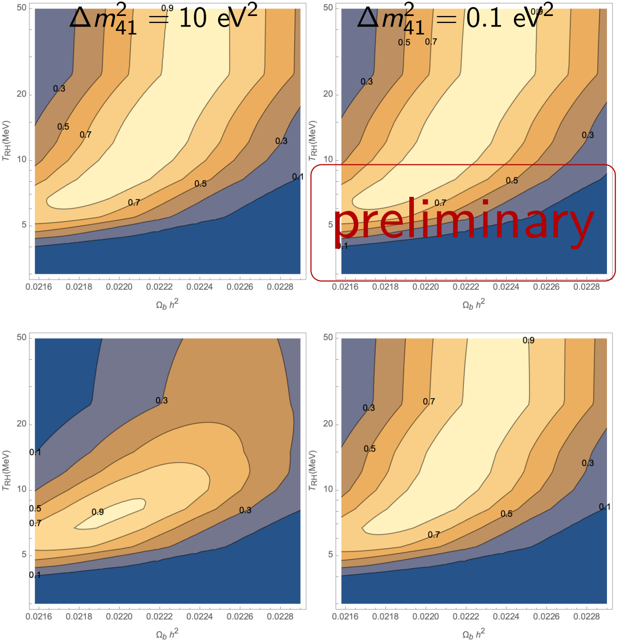
active,
sterile
contribution
to $N_{\rm eff}$

for low $T_{\rm rh}$, mixing parameters are irrelevant

for higher Δm^2_{41}, $T_{\rm rh}$ has more impact

BBN and low reheating

[in preparation]



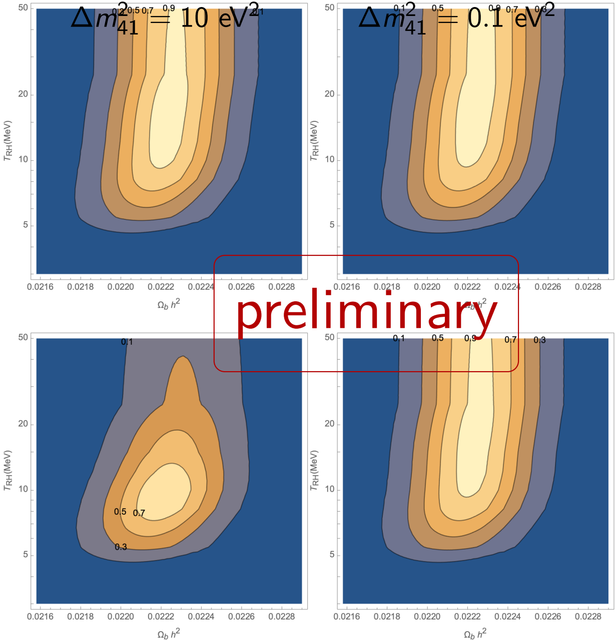
$|U_{e4}|^2 = 10^{-5}$

BBN (D+He) only

$|U_{e4}|^2 = 10^{-4}$

BBN and low reheating

[in preparation]



$\Delta m^2_{41} = 10\ \mathrm{eV}^2$

$\Delta m^2_{41} = 0.1\ \mathrm{eV}^2$

preliminary

$|U_{e4}|^2 = 10^{-5}$

BBN + Planck18

$|U_{e4}|^2 = 10^{-4}$

F

FELLINI

A few words on my FELLINI experience

Scientific production

- Secondments
    - 6 months at UTFSM, Santiago de Chile with prof. Aristizabal-Sierra (arxiv:2308.15531)
    - 6 months at IFIC, Valencia (Spain) with dr. Sergio Pastor (work on low-reheating scenarios in progress)
- Publications
    - 1 book chapter (arxiv:2306.15067, for the book "Hubble Constant Tension", Springer, expected in 2024)
    - 15 scientific papers on peer-reviewed journals (1 still under review)
    - 1 invited short review
- Dissemination
    - 12 talks at international workshops, conferences (including FELLINI general meetings)
    - 1 poster presentation at ICRC 2023
    - 3 lecture series at 3 international schools
    - 6+1 seminars at institutions in 5 different countries
- Public codes
    - FortEPiaNO (neutrino decoupling)
    - PArthENoPE v3 (BBN)

Other Activities

- Training
  - English course at UniTO
  - Machine learning hackathon at INFN
  - Machine learning course on Coursera
  - Course on project management at INFN
  - Course on HTCondor for system admins at INFN
  - Course “RESPECT” on inclusivity at INFN
- Outreach
  - Participation at European Researchers’ Nights
  - Participation with INFN Turin at Salone Internazionale del Libro (Turin)
- Funding applications based on my FELLINI project
  - ERC Starting Grant (scored B)
  - CIDEGENT (Spain)
  - La Caixa Foundation “Junior Leader” COFUND Fellowship (Spain/Portugal) – started on 1/10 at IFT (CSIC-UAM, Madrid)
  - FIS (in preparation)

Z

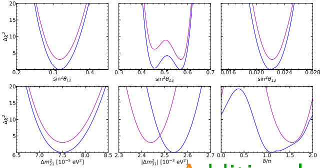
Conclusions

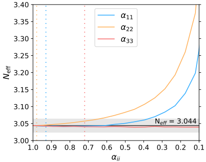
almost there!

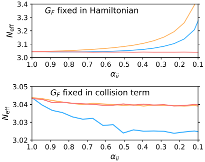
What do we know about neutrinos?

U

U: mixing matrix

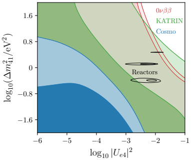


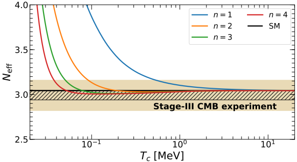


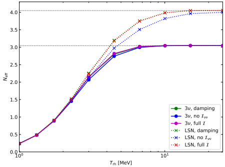


A

Additional particles/interactions

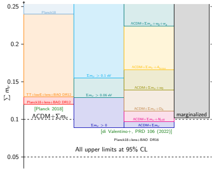


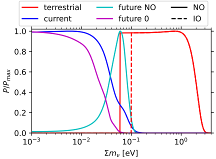


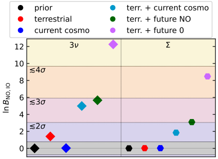


M

Masses and mass ordering







What do we know about neutrinos?

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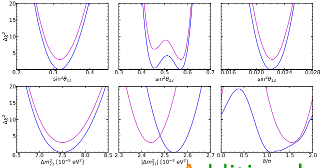

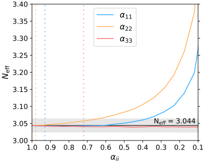

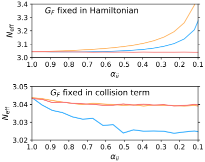

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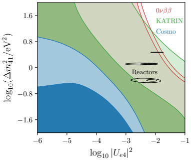

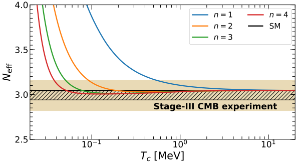

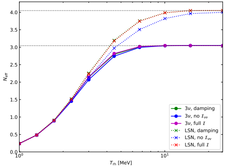

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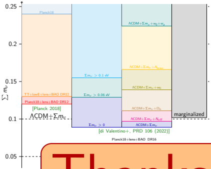

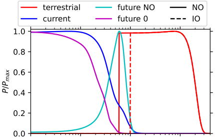

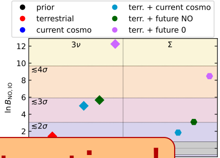

Thanks for your attention!